

HIGH-RESOLUTION ABSORPTION SPECTROSCOPY OF MULTI-PHASE, HIGH-METALLICITY GAS ASSOCIATED WITH THE LUMINOUS QUASAR HE 0226-4110 *

RAJIB GANGULY^{1,2}, KENNETH R. SEMBACH¹, TODD M. TRIPP³, BLAIR D. SAVAGE⁴, BART P. WAKKER⁴

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ABSTRACT

We present FUSE and HST/STIS observations of the absorption line system near the emission redshift of the radio-quiet, X-ray bright quasar HE0226-4110 ($z = 0.495$, $V = 15.2$). The spectra cover the rest-frame wavelength range 610–1150 Å, and we detect a wide range of ionization species, including four adjacent stages of oxygen: O III, O IV, O V, and O VI. Strong transitions of O I and O II are covered in our spectra, but none are detected. The detection of multiple ionization stages of a common element (oxygen) reveals a striking change in gas kinematics with ionization. Comparison of the O VI $\lambda\lambda 1031, 1037$ apparent column density profiles reveals no evidence for partial coverage or unresolved saturated structure, although parts of the O VI $\lambda 1037$ line are blended with Galactic C IV absorption. In addition, several transitions (e.g. C III $\lambda 977$, H I Ly- β) show black saturation which also indicates no unocculted flux that may dilute the absorption profiles. O III is only detected in a narrow feature which is also traced by the H I and C III lines, suggesting that they arise in the same gas. Absorption at the same velocity is also present in other species (N IV, O IV-VI, S IV, and possibly Ne VIII), but the kinematics differ from the O III, implying production in separate gaseous phases. The combination of H I, O III, and C III column density measurements with limits on the amount of O II and O IV in this phase yield an estimate of both the photoionization parameter and metallicity: $[\text{O}/\text{H}] = +0.12^{+0.16}_{-0.03}$, $\log U = -2.29^{+0.02}_{-0.23}$. We discuss two possible locations for the gas in this associated absorption-line system: the narrow emission line region of the quasar, and the halo of the quasar host galaxy. An additional narrow component that is only detected in O VI appears 58 km s⁻¹ redward of the O III-bearing gas. The narrow width of this component rules out collision-based ionization processes, and we constrain the ionization parameter of this component to $-0.35 \lesssim \log U \lesssim 0.02$. Additional structure is detected in the associated absorber in the form of two broad components. The components flank the narrow O III component in velocity and are detected in N IV, O IV-VI, and Ne VIII. The kinematics of O V and O VI in both components trace each other. The kinematics of N IV, O IV, and Ne VIII differ from O V and O VI and it is not clear how many phases or what ionization mechanism produces the gas in these broad components. A consideration of collisional ionization and photoionization equilibrium models, as well as radiative cooling and shock ionization models, leads us to conclude that Ne VIII must arise in a separate high-ionization phase.

Subject headings: quasars: absorption lines — quasars: emission lines — quasars: individual (HE 0226-4110) — ultraviolet: ISM

1. INTRODUCTION

Early after the discovery of quasars, it was recognized that their extreme luminosities, non-thermal spectra, variability, and compact sizes were readily achieved through the release of gravitational energy of matter accreting onto supermassive black holes (e.g., Greenstein & Schmidt 1964; Schmidt 1963). The remarkable result from the previous decade of work

that nearly all galaxies host central supermassive black holes (e.g., Tremaine et al. 2002; Gebhardt et al. 2000; Ferrarese & Merritt 2000; Magorrian et al. 1998) seems to suggest that any galaxy has the potential to host a quasar (or other form of active nucleus). In tandem with the observed fact that the quasar luminosity function peaks at $z \sim 2$ (e.g., Boyle et al. 2000) and falls off thereafter suggests that the quasar phenomenon is possibly a phase of galaxy evolution. Consequently, in order to trace the history of gas in the universe, it is important to understand how accretion disks around supermassive black holes are fueled, the dynamics and geometry of the accreting/outflowing gas, and how some fraction of the gas is returned to the IGM (i.e. feedback). A complete description of how gas falls into galaxies and is subsequently processed should include an understanding of the physical origin of absorption line systems that appear to cluster around quasars (e.g., Weymann et al. 1979; Foltz et al. 1986, 1988, hereafter “associated” absorption lines, or AALs⁵). Are there dynamical/kinematic rela-

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¹ Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218

² Department of Physics & Astronomy, University of Wyoming (Dept. 3905), 1000 East University Avenue, Laramie, WY 82071

³ Department of Astronomy, University of Massachusetts, Amherst, MA 01003

⁴ Department of Astronomy, University of Wisconsin-Madison, 475 N. Charter Street, Madison, WI 53706

⁵ Following Foltz et al. (1986), we refer to any narrow absorption line that arises within 5,000 km s⁻¹ as an AAL. We distinguish

tionships between the gas and those regions? Where is the gas located? What are the physical properties of the gas? Are AALs just weather around quasars or do they probe physically interesting regions around the black hole like the broad/narrow emission line region? Are AALs an integral feature of the quasar phenomenon?

The current paradigm of quasars holds that the disk of matter accreting onto the central supermassive black hole that powers quasars is also the origin of a wind (e.g., Elvis 2000; Proga, Stone, & Kallman 2000; Murray et al. 1995). This wind is thought to result from gas that is lifted off the accretion disk via local radiation pressure and then radially driven away from the black hole via scattering of ultraviolet lines. This scenario has been quite successful at explaining the existence and shape of both broad absorption lines (BALs) and the high-ionization broad emission lines. Low-ionization and high-excitation lines are more complicated as these may have a contribution from the outer regions of the accretion-disk itself (e.g., Eracleous & Halpern 2004). There are indications, including temporal variability of profiles and partial occultation of the absorber over the background quasar, that some fraction of narrow absorption-line systems also probe this outflowing gas (e.g., Ganguly et al. 2003; Misawa et al. 2003; Yuan et al. 2002).

In sight lines to high-redshift QSOs ($z \sim 2.5$) that are selected based on radio brightness, as many as 30% of C IV-selected systems with $z_{\text{abs}} \ll z_{\text{em}}$, commonly assumed to arise from the halos of intervening galaxies, may be intrinsic to the QSO (Richards et al. 1999; Richards 2001). This frequency was estimated by examining differences in the distribution of C IV-selected systems in apparent ejection velocity (over the range $+75,000 \geq v_{\text{ej}} \geq -6,000 \text{ km s}^{-1}$) as a function of quasar radio loudness and spectral index. If C IV-selected systems are unrelated to the background quasar (as in the case of halos from cosmologically-distributed interloping galaxies), then no variation is expected. In the apparent ejection velocity range $+6,000$ to $+75,000 \text{ km s}^{-1}$ (i.e., at large separations from the quasar redshift), Richards (2001) found a statistical excess of C IV-systems toward radio-quiet quasars compared to radio-loud quasars, and in flat-spectrum quasars compared to steep-spectrum quasars.

This estimate of the frequency of high-ejection velocity systems is surprisingly high given that, at lower redshifts, many studies have shown that absorption systems with $z_{\text{abs}} \ll z_{\text{em}}$ are strongly correlated with foreground galaxies, indicating that these low- z absorbers are *not* intrinsic systems with high ejection velocities. It is not yet clear how these observations can be reconciled. Studies to understand the populations of absorbing galaxies at high-redshift (i.e., $z_{\text{abs}} \gtrsim 1.5$) are not yet feasible with current technology. Likewise, it is possible that the selection of optically-bright low-redshift quasars in follow-up studies of absorbing galaxy populations has an effect on the velocity distribution of intrinsic systems such that they are not observed.

In a heterogeneous study of low-redshift quasars, Wise et al. (2004) have found that at least 20% of low- z AALs (i.e., systems appearing within 5000 km s^{-1} of the

quasar redshift) show time variability and therefore are likely to be intrinsic. While intrinsic absorbers that lie at large velocity separations from the quasar likely originate in some form of high-velocity outflow, there are additional possible locations for AALs: ablation off an obscuring torus (Krolik & Kriss 2001), the narrow emission line region, the quasar host galaxy (e.g., Baker et al. 2002), satellite dwarfs, or luminous companion galaxies in the quasar’s host cluster.

There has been considerable statistical work on the relationship between associated absorbers and various quasar properties. Foltz et al. (1986, 1988) established that intermediate redshift AALs ($1.43 \lesssim z \lesssim 1.94$) with large C IV equivalent widths [$W_{\text{rest}}(\text{C IV}) \gtrsim 1.5 \text{ \AA}$] preferentially appear toward steep-spectrum radio-loud quasars, in stark contrast with BALs, which prefer more radio-quiet quasars (e.g., Becker et al. 2000). While the dichotomy in radio-loudness in quasars has apparently been eliminated through deep radio imaging by the FIRST survey (Becker et al. 1995), there remains an unexplained discrepancy between quasar radio-loudness and the presence of intrinsic absorption. Ganguly et al. (2001) reported an apparent dearth of strong AALs among the lower redshift ($z < 1$) quasars observed with the *Hubble Space Telescope* (HST) for the HST Quasar Absorption Line Key Project, suggesting possible evolution in the frequency of strong AALs. While the absence of strong AALs in that sample might be explained partially as a bias against bluer quasars, there is still evidence for evolution from larger and well-selected quasar samples (Brandt et al. 2000; Vestergaard 2003).

As noted by several authors (e.g., Mathur et al. 1995; Brandt et al. 2000), there is apparently a connection between C IV-selected AALs, and absorption at soft X-ray wavelengths (so-called “warm absorbers”). The precise nature of this connection is perplexing given the drastically different ionization potentials of species (e.g., O VII, O VIII) that exist in warm absorbers. A number of campaigns have been carried out employing simultaneous HST Space Telescope Imaging Spectrograph (HST/STIS), *Far Ultraviolet Spectroscopic Explorer* (FUSE), and *Chandra X-ray Observatory*/XMM-Newton spectroscopy of Seyfert galaxies, the lower-luminosity kin of quasars, such as NGC 3783 (Netzer et al. 2003; Gabel et al. 2003a,b; Kaspi et al. 2002), Mrk 279 (Scott et al. 2004b), and NGC 5548 (Crenshaw et al. 2003). The results of these studies of Seyfert galaxies in regards to the general relationship between the X-ray and UV absorbing gas is inconclusive. In some cases, the absorbers can be explained by the same phase of gas; in other cases, a more complex, multi-phase absorber is required. For the more luminous quasar 3C 351, Yuan et al. (2002) have analyzed high resolution UV and X-ray spectra and find that complex, multi-phase gas is required to explain the components in the AAL.

Like the previously mentioned campaigns, it is important to temper statistical relationships between absorbers and quasar properties with case studies of individual objects. Since the demographics of the absorbers are clearly eclectic, only case-by-case classifications to identify subsamples can hope to recover true, undiluted physical relationships. Using HST/STIS and FUSE ob-

this purely observational classification from absorption that is truly related (i.e., intrinsic) to the quasar central engine or host galaxy.

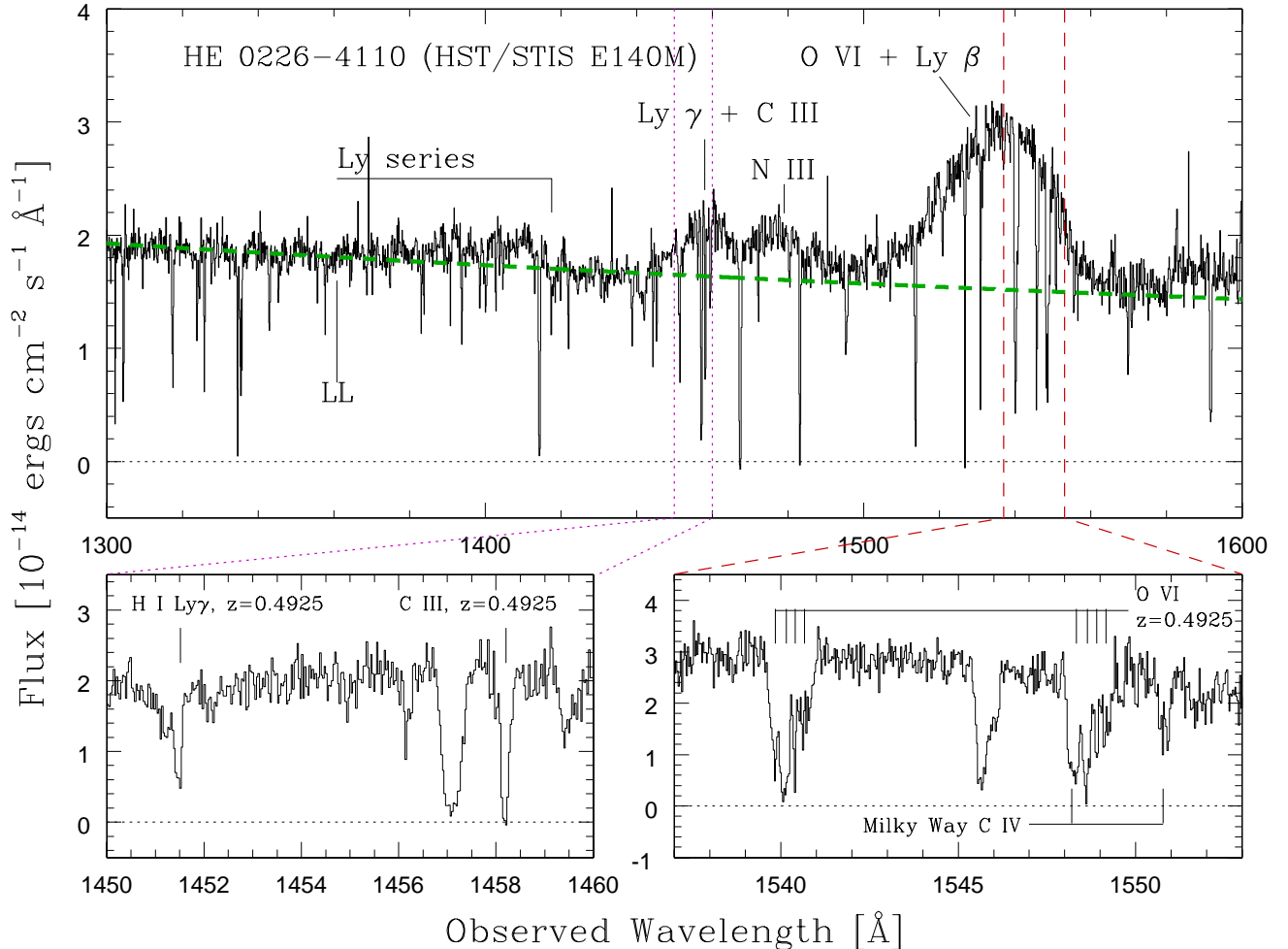


FIG. 1.— Portions of the STIS E140M echelle spectrum of HE0226-4110. In order to improve the signal-to-noise ratio and thereby show the continuum and emission-line shapes more clearly, the upper panel is binned to a sampling of $\sim 20 \text{ km s}^{-1}$. In the upper panel, the thick dashed line shows the power law fitted to the continuum for the purposes of the photoionization models presented in the text (see §3.1.4.1). The O VI+Ly β , N III, Ly γ +C III, and higher order Lyman series emission lines of the QSO are marked. An ‘LL’ marks the location of the Lyman limit for the associated absorption-line system. The lower panels show selected associated absorption lines at $z_{\text{abs}} = 0.4925$ including H I Ly γ and C III $\lambda 977.020$ (lower left panel) and the O VI $\lambda\lambda 1031.926, 1037.617$ doublet (lower right panel). The lower panels are binned to $\sim 7 \text{ km s}^{-1}$ pixels. Note that the O VI $\lambda 1037.617$ line is blended with one of the lines of the Galactic C IV $\lambda\lambda 1548.204, 1550.781$ doublet, which is also marked.

servations of low-redshift quasars that cover the O VI $\lambda\lambda 1031.927, 1037.617$ doublet, not only can we bridge the gap between C IV-selected AALs and soft X-ray “warm absorbers,” but we can also take advantage of the high spectral resolution to understand the physical properties and origin of the gas. Kriss (2002) presents a summary of FUSE observations of bright AGN at low-redshift (predominantly Seyfert galaxies). These observations survey the O VI doublet in AGN out to $z \sim 0.15$. In a companion paper, we extend this survey out to $z \sim 0.5$ (Ganguly et al. 2005, in prep.).

In this paper, we report the discovery of a remarkable absorption line system ($z_{\text{abs}} = 0.4925$) near the redshift of the radio-quiet, X-ray bright quasar HE0226-4110 ($z_{\text{em}} = 0.495, V = 15.2$; Reimers et al. 1996). In §2, we discuss the details of our HST/STIS and FUSE observations. In §3, we investigate the properties of the quasar to provide a context for understanding the asso-

ciated absorption presented and analyzed in §4. Finally, in §5, we summarize our results and discuss the implications of our findings.

2. DATA

2.1. STIS Observations

Our HST/STIS (Woodgate et al. 1998; Kimble et al. 1998) observations were carried out using the $0''.2 \times 0''.06$ slit, the E140M grating, and the NUV MAMA detector between 25 December 2002 and 1 January 2003. The total integration time of the observations was 43.8 ksec. This configuration provides a spectral resolution of about 6.5 km s^{-1} (FWHM) with a sampling of 2–3 pixels per resolution element, and yields semi-continuous spectra over the observed wavelength range 1149–1729.5 Å. The data were reduced and calibrated as described by Tripp et al. (2001). For further details on the processing of the datasets used in this work for the HE0226-4110

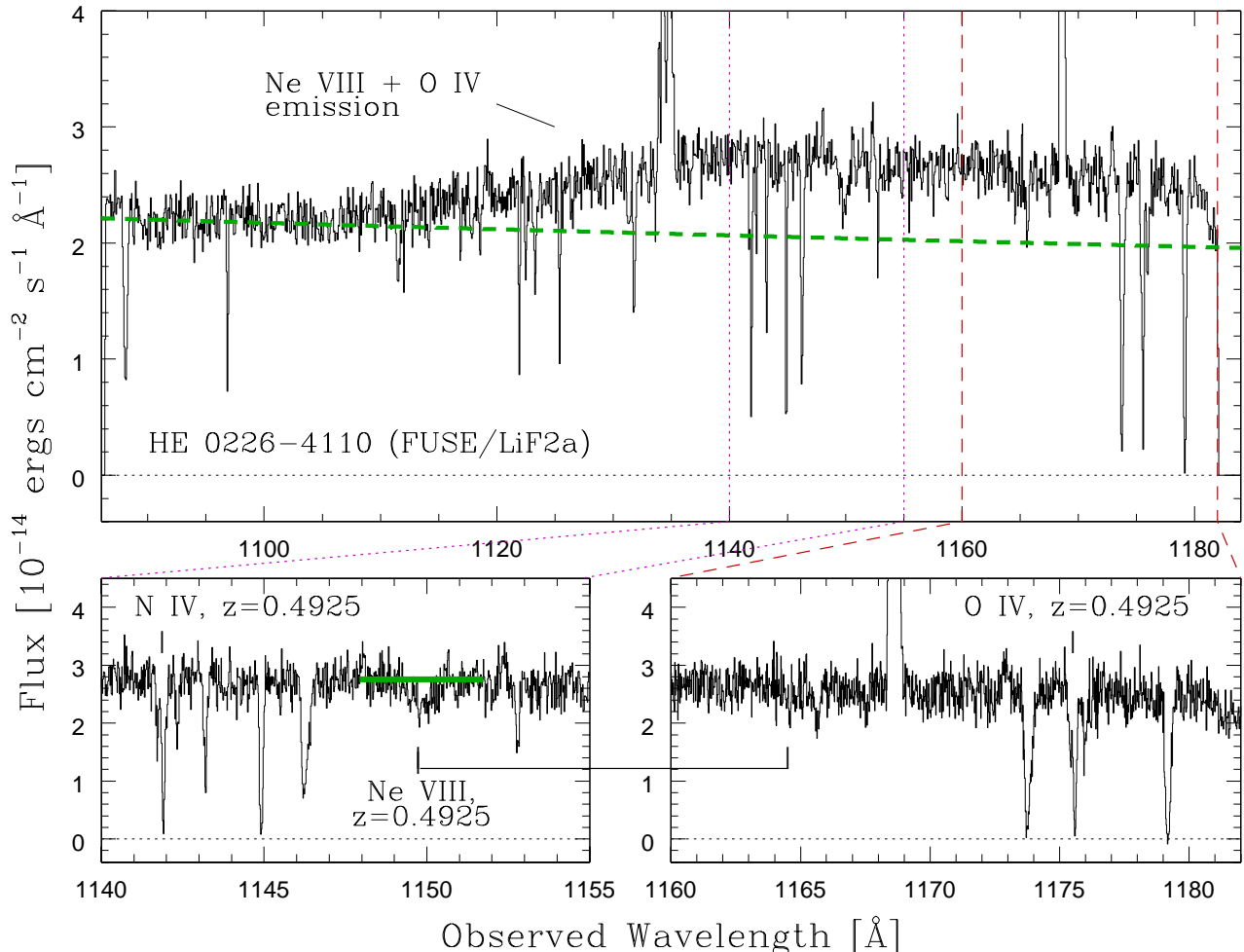


FIG. 2.— Portions of the FUSE LiF2a spectrum of HE0226-4110. In order to improve the signal-to-noise ratio and thereby show the continuum and emission-line shapes more clearly, the upper panel is binned to a sampling of $\sim 20 \text{ km s}^{-1}$. In the upper panel, the thick dashed line shows the power law fitted to the continuum for the purposes of the photoionization models presented in the text (see §§3.1, 4.1), and the Ne VIII+O IV emission line of the QSO is marked. The lower panels show selected associated absorption lines at $z_{\text{abs}} = 0.4925$ including N IV $\lambda 765.148$ (lower left panel), O IV $\lambda 787.711$ (lower right panel), and the Ne VIII $\lambda\lambda 770.409, 780.324$ doublet (both lower panels). In the lower left panel, our local Legendre-polynomial fit to the continuum around the Ne VIII $\lambda 770.409$ line is shown as a thick line. The Ne VIII $\lambda 780.324$, which would nominally appear in the lower right panel, is not detected. The lower panels are binned to a sampling of $\sim 6 \text{ km s}^{-1}$.

sight line, see Fox et al. (2005) and Savage et al. (2005). According to Kim Quijano et al. (2003), the zero-point heliocentric velocity uncertainty in the calibration is about 0.2–0.5 pixels, or $\sim 0.6 - 1.5 \text{ km s}^{-1}$. The flux calibrations are good to 8% (or roughly $1 - 2 \times 10^{-15} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Å}^{-1}$ for these data). The integrations yielded a final signal-to-noise per resolution element (after reduction, calibration, and co-addition of spectra) of ~ 11 at 1300 Å and ~ 8.5 at 1500 Å .

Selected portions of the STIS echelle spectrum of HE0226-4110 are shown in Figure 1. The upper panel of Figure 1 shows a broad section of the spectrum including the O VI+Ly β broad emission line of the QSO as well as several weaker emission features. We also show in the upper panel a power law fitted to the QSO continuum (see §3.1) which we will use for photoionization modeling in §4.1. The spectrum is rich in Galactic and extragalactic absorption lines. Extragalactic interven-

ing absorption systems (i.e., with $z_{\text{abs}} \ll z_{\text{em}}$) in the spectrum of HE0226-4110 are analyzed and discussed in Savage et al. (2005) and Lehner et al. (2006), and Galactic absorption lines and high-velocity clouds toward this QSO are discussed by Collins et al. (2005) and Fox et al. (2005).

Examples of the associated absorption lines of interest in this paper that are redshifted into the STIS-E140M band are shown in the lower panels of Figure 1 including the multicomponent O VI $\lambda\lambda 1031.926, 1037.617$ doublet at $z_{\text{abs}} = 0.4925$ (lower right panel) and the H I Ly γ and C III $\lambda 977.020$ lines (lower left panel). It is immediately evident that this is a complex, multiphase absorber: four components are present in the O VI lines spread over a velocity range of 250 km s^{-1} , but only a single component is seen in tracers of lower-ionization gas such as C III and Ly γ (the weak component blueward of the Ly γ line shown in Figure 1 is an unrelated Ly α line at a different

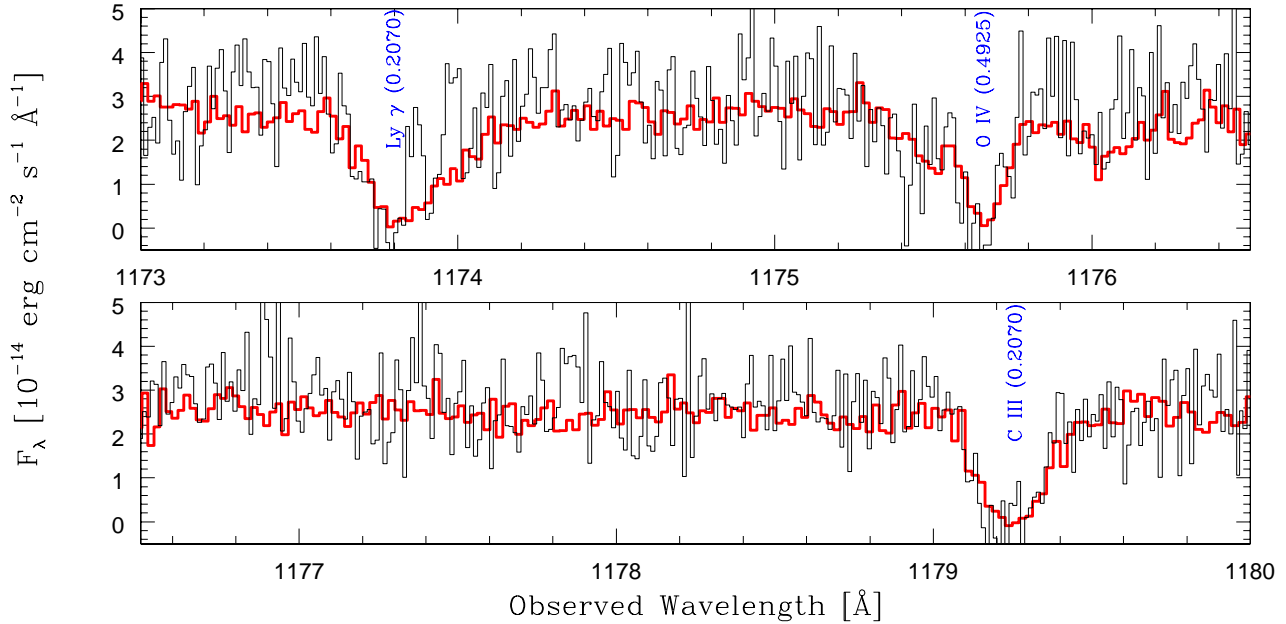


FIG. 3.— We show a portion of the overlapping region between the FUSE LiF2A (thick histogram) and STIS E140M (thin histogram) spectra. The plotted region contains associated absorption from O IV, as well as IGM absorption from Lyman γ and C III.

redshift). We also see from this figure that the associated O VI $\lambda 1037.617$ line at $z_{\text{abs}} = 0.4925$ is blended with Milky Way C IV $\lambda 1548.204$. This blend is not a serious problem because the O VI $\lambda 1031.926$ line is free of blending, and the Galactic C IV $\lambda 1550.781$ transition is also unblended, which enables an assessment of the degree of contamination from C IV $\lambda 1548.204$.

2.2. FUSE Observations

Our FUSE observations were carried out using the LWRS aperture between 12 December 2000 and 21 October 2003 for a total integration time of 208.9 ksec. In this mode, the FUSE satellite delivers spectra with velocity resolution in the range $20\text{--}30\text{ km s}^{-1}$, with 10 pixels per resolution element. For further details about FUSE and its on-orbit performance, see Moos et al. (2000) and Sahnou et al. (2000). The data were reduced using version 2.4.0 of the CALFUSE pipeline; see Fox et al. (2005) and Savage et al. (2005) for further details regarding subsequent wavelength calibration steps and combination of sub-integrations. (We note that over the 34 month span of the observations, the quasar continuum level increased by about 25%.) In this study, we use primarily the LiF2A and SiC2A detector segments that cover the wavelength ranges $1086.3\text{--}1182.0\text{ \AA}$ and $916.5\text{--}1006\text{ \AA}$, respectively. In addition, the FUSE data are rebinned to a sampling of about 3-4 bins per resolution element. The CALFUSE pipeline and subsequent processing provides fully calibrated (flux and wavelength) spectra that are accurate to 5 km s^{-1} in wavelength and about 10% ($1 - 2 \times 10^{-15}\text{ erg cm}^{-2}\text{ s}^{-1}\text{ \AA}^{-1}$) in flux. [Note that the FUSE spectra were aligned with the STIS spectra using Galactic and intervening absorption lines; see §2.3.] The long integration time on this bright quasar provided signal-to-noise of 23 per resolution element at 1150 \AA (in the LiF2A detector segment) and 13 per resolution element at 950 \AA

(in the SiC2A detector segment).

Selected portions of the FUSE LiF2A spectrum are shown in Figure 2. The upper panel of the figure shows the entire portion of the FUSE spectrum that is covered by the LiF2A detector segment. For clarity, the spectra in this panel are shown with a sampling of $\sim 20\text{ km s}^{-1}$, or 1 bin per FUSE resolution element in this region of the spectrum. The underlying QSO continuum from the power law fit to the STIS spectrum is also shown; the contribution of the Ne VIII and O IV broad emission lines is readily apparent.

In the bottom panels, we show examples of the associated absorption lines that are clearly detected in this spectrum. In these panels, the spectra are shown with a sampling of $\sim 6\text{ km s}^{-1}$. The lower left panel shows the portion of the spectrum covering the associated absorption from the N IV $\lambda 765.148$ and Ne VIII $\lambda 770.409$ lines. In this panel, we also overlay our Legendre-polynomial fit to the local continuum around the Ne VIII $\lambda 770.409$ line to show that the detection of the feature is robust. The line, which appears at an observed wavelength of $\sim 1150\text{ \AA}$, is also readily apparent in the top panel. In the bottom right panel, we show another region of the FUSE LiF2A spectrum that covers the Ne VIII $\lambda 780.324$ and O IV $\lambda 787.711$ lines. The Ne VIII $\lambda 780.324$ line is not detected in our spectrum down to a 3σ equivalent width limit of 103 m\AA . The Ne VIII $\lambda 770.409$ line has an equivalent width of $140 \pm 22\text{ m\AA}$, so the non-detection of the weaker line is consistent with a physically meaningful doublet ratio (i.e., $1 \leq W_{\lambda}(\lambda 770.409)/W_{\lambda}(\lambda 780.324) \leq 2$). We conclude that the detection of the absorption feature at $\sim 1150\text{ \AA}$ and its identification with the Ne VIII $\lambda 770.409$ line is reliable.

2.3. A Remark on the Wavelength Cross-Calibration

Of particular importance to this study, the FUSE and STIS data were aligned to a common (heliocentric) wave-

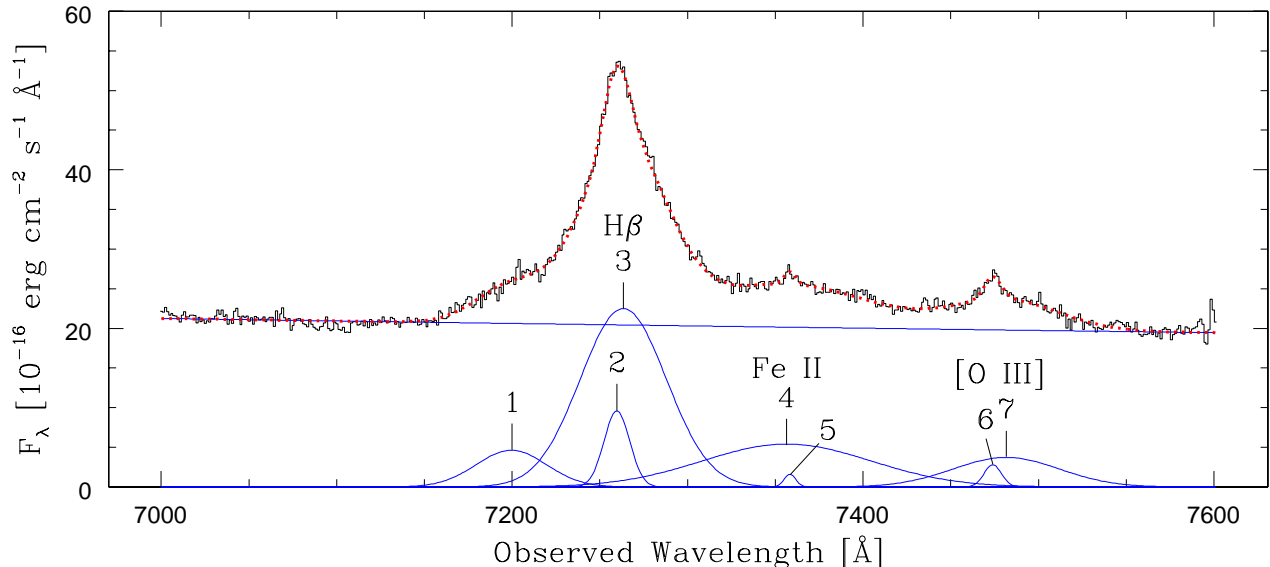


FIG. 4.— We show a portion of the Grupe et al. (1999) optical spectrum of HE 0226-4110 (histogram) which covers the H β and [O III] λ 5007 lines. The smooth curves show the contributions to the best-fit from the underlying power-law continuum and each emission line component. The dotted curve shows the total best-fit spectrum. Each emission line component is coded with a number that appears over the component that is used to cross-reference with the fit parameters listed in Table 1. The locations of the H β , Fe II, and [O III] emission lines are also labelled.

length scale using the Galactic C II $\lambda\lambda$ 1334.532, 1036.337 and Fe II $\lambda\lambda$ 1144.938, 1608.451 lines, as well as the Lyman γ and C III λ 977.020 IGM lines of the $z_{\text{abs}} = 0.2070$ absorber. The latter two lines appear in the wavelength region 1149–1182 Å where the FUSE and STIS spectra overlap. A portion of the overlapping region is shown in Figure 3. The shift required to align the FUSE spectra to the heliocentric velocity scale provided by the STIS calibration was 3–5 km s $^{-1}$, which is within the residual zero-point uncertainty reported by Fox et al. (2005). The alignment of these spectra is important to this study, because there is an apparent ionization-dependent velocity shift in the AAL at $z_{\text{abs}} = 0.4925$ between the O III λ 832.927 line and the O IV λ 787.711 line (see §4.1). The former line appears in the STIS spectra, while the latter line appears in the overlapping region and is flanked by two IGM lines from the same system (at $z_{\text{abs}} = 0.2070$). The alignment of the IGM features gives us confidence in the magnitude of the shift, and, consequently, in the reality of the apparent velocity shift in the associated absorber.

2.4. Comments on Ne VIII

The detection of the high-ionization Ne VIII λ 770.409 line is fortuitous as it provides a further bridge between lower-ionization associated C IV absorption and the higher-ionization X-ray absorption by O VII and O VIII (e.g., Mathur et al. 1995; Brandt et al. 2000). To our knowledge, Ne VIII has only been detected in associated narrow absorption in three other quasars, UM 675 (Hamann et al. 1995), 3C 288.1 (Hamann et al. 2000), HDF5-QSO J2233 – 606 (Petitjean & Srianand 1999), and possibly HS 1700+6416 (Petitjean et al. 1996). Previous detections of Ne VIII absorption in BALQSOs include PG 0946+301 (Arav et al. 1999), SBS 1542+541 (Telfer et al. 1998), and Q 0226-1024 (Korista et al. 1992). We note also that the FUSE *composite* spectrum

of Scott et al. (2004a) shows a peculiar absorption feature just blueward of the Ne VIII+O IV emission at a rest wavelength of ~ 730 Å. [As noted by Scott et al. (2004a), this feature was also present, though at a much weaker level, in the rest-frame extreme ultraviolet HST composite from Telfer et al. (2002).] The HE 0226-4110 FUSE spectrum presented here was used in that composite, but the associated absorber does not contribute to that feature.

3. QUASAR PROPERTIES

Before proceeding with an analysis of the associated absorption at $z_{\text{abs}} = 0.4925$, we first discuss the general UV/optical, radio, and X-ray properties of the quasar, as this is important to understanding the environment of the absorbing gas.

3.1. UV/Optical

The redshift of the quasar as measured by the peak of the Mg II emission line (Reimers et al. 1996) is $z_{\text{em}} = 0.495 \pm 0.001$. In general, it is more favorable in the study of associated absorbers to use redshifts determined from narrow emission lines (as these can be measured to higher precision) or from high-excitation lines which are thought to arise from the quasar accretion disk (not the outflow), and may more accurately reflect the systemic redshift. To this end, we have analyzed the optical spectrum obtained by Grupe et al. (1999) using the ESO Faint Object Spectrograph and Camera (≈ 5 Å FWHM resolution with 1.2 Å/pixel sampling). The spectrum covers the H β and [O III] λ 5007 emission lines and is shown in Figure 4. (At the redshifted position of the [O III] λ 5007 line, the resolution of the spectrum is ≈ 200 km s $^{-1}$ with a sampling of 48 km s $^{-1}$.) Since the H β and [O III] λ 5007 emission lines lie on top of an Fe II multiplet, we model the spectrum with an underlying power-law continuum

TABLE 1. PARAMETERS OF FIT TO OPTICAL SPECTRUM

Comp.	w_i^a	λ_i (Å)	FWHM _i ^b (km s ⁻¹)	Ident.	z_{em}	$\log \lambda L_\lambda^c$ (erg s ⁻¹)
1	0.22 ± 0.02	7199.6 ± 2.9	1968 ± 178	UID	...	43.33 ± 0.05
2	0.45 ± 0.07	7259.7 ± 1.0	710 ± 124	H β	0.4929 ± 0.0001	43.21 ± 0.10
3	1.06 ± 0.06	7263.4 ± 1.0	2376 ± 117	H β	0.4937 ± 0.0001	44.10 ± 0.03
4	0.25 ± 0.01	7356.5 ± 3.1	4610 ± 414	Fe II	...	43.78 ± 0.04
5	0.07 ± 0.02	7358.4 ± 0.8	319 ± 85	UID ^d	...	42.09 ± 0.16
6	0.13 ± 0.02	7474.0 ± 0.8	466 ± 86	[O III] $\lambda 5007$	0.4928 ± 0.0001	42.51 ± 0.11
7	0.18 ± 0.01	7481.6 ± 2.0	2861 ± 191	[O III] $\lambda 5007$	0.4942 ± 0.0002	43.42 ± 0.04

NOTE. — Uncertainties are quoted at 1σ confidence. The best-fit flux normalization at 7000 Å is $F_{\lambda_o} = (2.125 \pm 0.05) \times 10^{-15} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Å}^{-1}$. The best-fit power-law index was $\beta = -1.07 \pm 0.09$ ($F_\lambda \propto \lambda^\beta$). See §3 for explanations of the parameters.

^aEmission line strengths are quoted relative to best-fit flux normalization at 7000 Å .

^bThe width of the emission line components is quoted as the full-width at half-maximum intensity ($2.35\sigma_i$).

^cThe line luminosity for each component was computed assuming isotropic radiation and a luminosity distance of 2774.9 Mpc ($z = 0.493$ for a $\Omega_\Lambda = 0.73$, $\Omega_m = 0.27$ cosmology).

^dOne possible identification of this emission line component is [O III] $\lambda 4959$. However, this is not consistent with either the centroids or strengths of the listed [O III] $\lambda 5007$ components. There are potential systematic effects resulting from an improper accounting of the underlying Fe II emission, but such an accounting is beyond the scope of this work.

with Gaussian emission line components:

$$F_\lambda = F_{\lambda_o} \left[\left(\frac{\lambda}{\lambda_o} \right)^\beta + \sum_{i=1}^m w_i e^{-\frac{1}{2} \left(\frac{\lambda - \lambda_i}{\sigma_i} \right)^2} \right], \quad (1)$$

where λ_o is an arbitrary reference wavelength, F_{λ_o} is the normalizing continuum flux at that wavelength, β is the continuum power-law index, m is the number of emission line components, and w_i , λ_i , and σ_i are the relative strength, central wavelength, and width of each emission line component, respectively. We used a Marquardt-Levenburg least-squares algorithm (Press et al. 1992) to determine the best-fit for a given number of emission components, and the F-test to choose the optimal number of statistically significant components. In our fitting algorithm, we use the entire spectrum and allow both power-law and emission line parameters to vary simultaneously to minimize the χ^2 (defined in the usual manner). The errors on all parameters are given by the diagonal elements of the covariance matrix (the inverse of the χ^2 curvature matrix). In Figure 4 we also show the best-fit model, as well as the separate contributions from the power-law continuum and each emission-line component. The parameters of the fit are listed in Table 1. The fit appears reasonable, especially in regions far from the emission lines, so we are confident in the parameters of the underlying power-law.

Both the H β and [O III] $\lambda 5007$ emission lines are well-described by the superposition of a broad component (FWHM = $2861 \pm 191 \text{ km s}^{-1}$) at $z_{\text{em}} = 0.4938 \pm 0.0001$ and a narrow component (FWHM = $966 \pm 86 \text{ km s}^{-1}$) at $z_{\text{em}} = 0.4928 \pm 0.0001$. Three additional components are required to yield a good fit to the data. We identify one component with the Fe II $\lambda 4924$ multiplet. Since our goal is only to measure the redshift and intensity of the [O III] $\lambda 5007$ line (to compare with the associated absorption in the O III $\lambda 832.927$ line), we do not attempt to further model the Fe II emission.

In addition to the fit of the optical continuum and emission lines, we have also fit the STIS spectrum which includes the Lyman limit of the associated absorber. This is important for characterizing the ionizing radiation field assumed later in our photoionization models.

Our method for fitting the UV spectrum was identical to that of the optical spectrum (described above), and the underlying power-law is shown in Figure 1. The extrapolation of this power-law to the FUSE band is also shown in Figure 2. In both figures, the fit appears quite reasonable. The parameters of the power-law fit are: $F_{\lambda_o} = (1.74 \pm 0.01) \times 10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Å}^{-1}$, $\beta = -1.42 \pm 0.02$ at $\lambda_o = 1395.55 \text{ Å}$. We recast this fit into frequency units ($F_\nu \propto \nu^{\alpha_{\text{UV}}}$) and to the rest-frame of the quasar: $F_{\nu_o} = (5.00 \pm 0.03) \times 10^{-27} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Hz}^{-1}$, $\alpha_{\text{UV}} = -0.58 \pm 0.02$ at 1 Rydberg. As in Table 1, we adopt a luminosity distance of 2774.9 Mpc ($z = 0.493$ for a $\Omega_\Lambda = 0.73$, $\Omega_m = 0.27$ cosmology). Thus, the luminosity density of the quasar at 1 Rydberg is $\log L_\nu [\text{erg s}^{-1} \text{ Hz}^{-1}] = 30.65$.

3.2. Radio

Another important quasar property to consider in understanding this associated absorber is the radio-loudness (defined as the ratio of flux densities at 5 GHz and 2500 Å , R^*), since it appears that absorption which is clearly related to quasar outflows (i.e., broad absorption lines) is connected to the presence and strength of radio emission. The precise nature of this connection is not yet clear. Strong, narrow, associated C IV absorption is preferentially found in steep-spectrum, radio-loud ($R^* > 10$) quasars (e.g. Foltz et al. 1988; Vestergaard 2003), while broad absorption lines are predominantly found in radio-quiet quasars (e.g. Turnshek 1988; Weymann et al. 1991; Brotherton et al. 1998; Becker et al. 2000).

This quasar is not detected by the Parkes-MIT-NRAO survey (Griffith & Wright 1993) down to a 4.4σ flux density limit of 47.3 mJy at 4.85 GHz . This limit corresponds to a radio-loudness of $\log R^* < 1$, which is radio-quiet by the definition of Kellermann et al. (1989). The radio-loudness is an apparently important indicator of the soft X-ray spectral slope (Laor et al. 1997), as well as being one of a collection of properties (including the [O III] $\lambda 5007$ luminosity, and H β emission line full-width at half-maximum intensity) that are apparently related among optically-selected quasars (that is, Eigenvector 1 from the principal components analysis

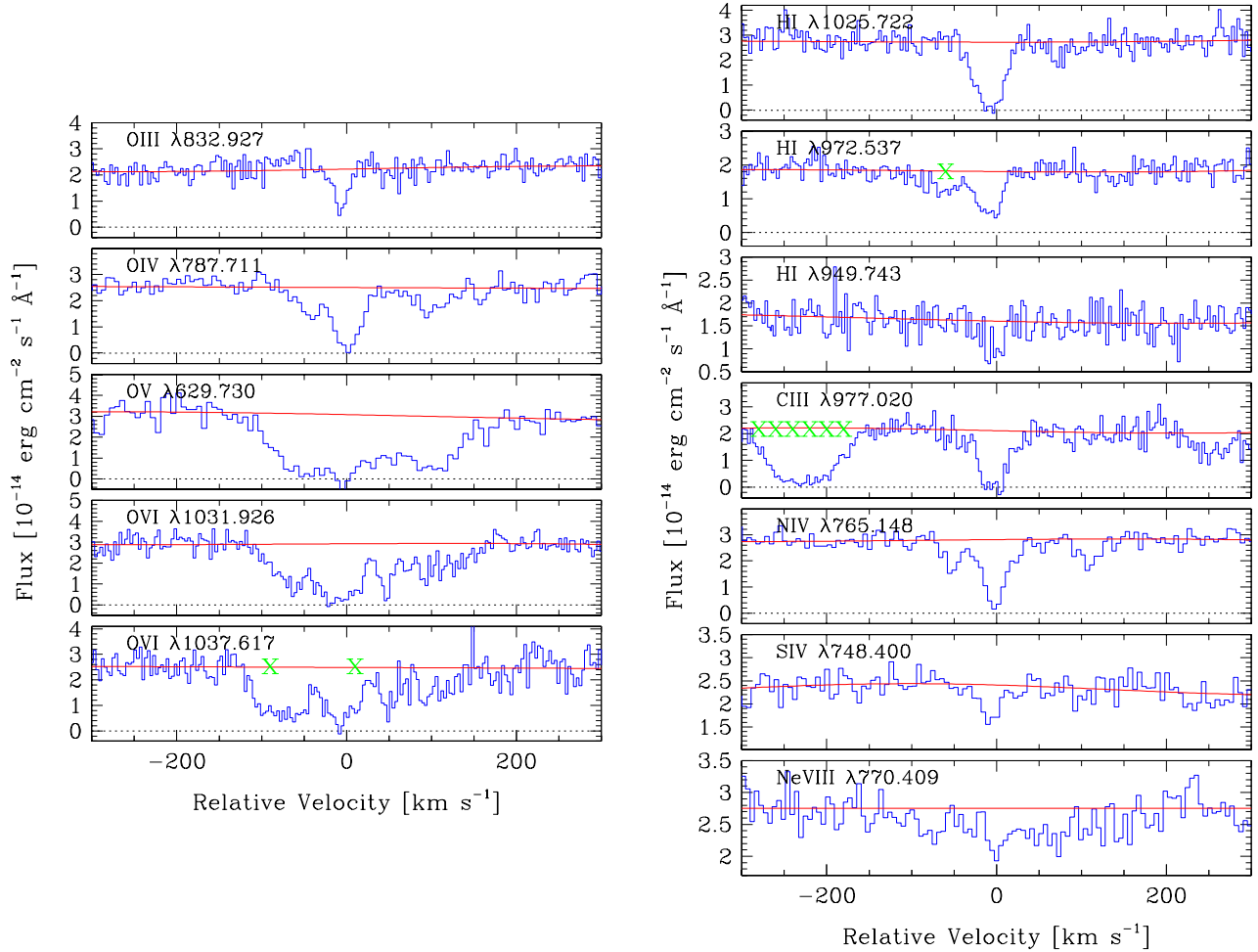


FIG. 5.— Plots of flux versus velocity (relative to $z_{\text{abs}} = 0.4925$) of ions detected in the FUSE and STIS spectra. Velocities where absorption is blended by a feature from an unrelated system are marked with an ‘x’. The smooth curves are our assessment of the local continuum using low-order Legendre polynomials. On the left, we plot all detected transitions from species of oxygen, which clearly show the ionization-dependent differences in the kinematics. On the right, we show transitions from all other detected species. Due to the shallowness of the Ne VIII $\lambda 770.409$ absorption, the profile and continuum fit are not well-represented in this expanded velocity range. We refer the reader to the bottom left panel of Figure 2 for a better view of this absorption profile and its continuum fit.

of Boroson & Green 1992). The underlying connection among these properties is generally interpreted as the accretion-rate of material onto the black hole relative to Eddington, M/M_{edd} . The radio-loudness of this quasar and its relatively weak [O III] $\lambda 5007$ emission indicate that the accretion-rate relative to Eddington is low in comparison with other quasars.

3.3. Soft X-ray

There seems to be a connection between absorption in ultraviolet transitions (in particular C IV $\lambda\lambda$ 1548.204, 1550.781) and “warm absorbers” in the soft X-rays (e.g., from O VII, and O VIII edges) from both studies of statistical coincidence (e.g., Brandt et al. 2000; Monier et al. 2001) and from analyses of specific sight lines (e.g. Scott et al. 2004b; Crenshaw et al. 2003; Gabel et al. 2003a,b; Mathur et al. 1995). Since the absorption in these two bands results from species with widely differing ionization potential, the precise causal connection is unclear. HE 0226-4110 was detected by the *ROSAT* All-Sky Survey (Voges et al. 1999, RASS) with a count rate of 0.56 ± 0.06 cps in the 0.2-2.0 keV bandpass

of the PSPC. Grupe et al. (1998) fit the RASS PSPC X-ray spectrum of HE 0226-4110 assuming the intrinsic spectrum is a single power-law and reported a best-fit spectral index of $\alpha_{\text{X}} = -2.1 \pm 0.2$ with a total absorbing column density of $N(\text{H}) = (1.78 \pm 1.2) \times 10^{20} \text{ cm}^{-2}$. The origin of this total absorbing column density can be attributed to Galactic absorption. [From a fit to damped Lyman α and β lines, Fox et al. (2005) find a Galactic H I column density of $1.2^{+0.4}_{-0.3} \times 10^{20} \text{ cm}^{-2}$. Similarly, Wakker et al. (2003) report an H I column density of $1.55 \pm 0.01 \times 10^{20} \text{ cm}^{-2}$ from a fit to 21 cm emission in this direction.] Grupe et al. (1998) find no evidence for a significant amount of the absorption in the soft- X-rays, and in particular from O VII and O VIII.

In further characterizing the ionizing radiation field, we compute the two-point 2 keV-to-2500 Å spectral index, α_{ox} , which provides a measure of the relative importance of the X-ray emission to the UV emission. Using our power-law fit to the STIS spectrum, and the Grupe et al. (1998) power-law fit to the *ROSAT* spectrum, we find $\alpha_{\text{ox}} = -1.46 \pm 0.04$, which is close to the average value for unabsorbed quasars ($\langle \alpha_{\text{ox}} \rangle \approx -1.5$; Brandt et al. 2000).

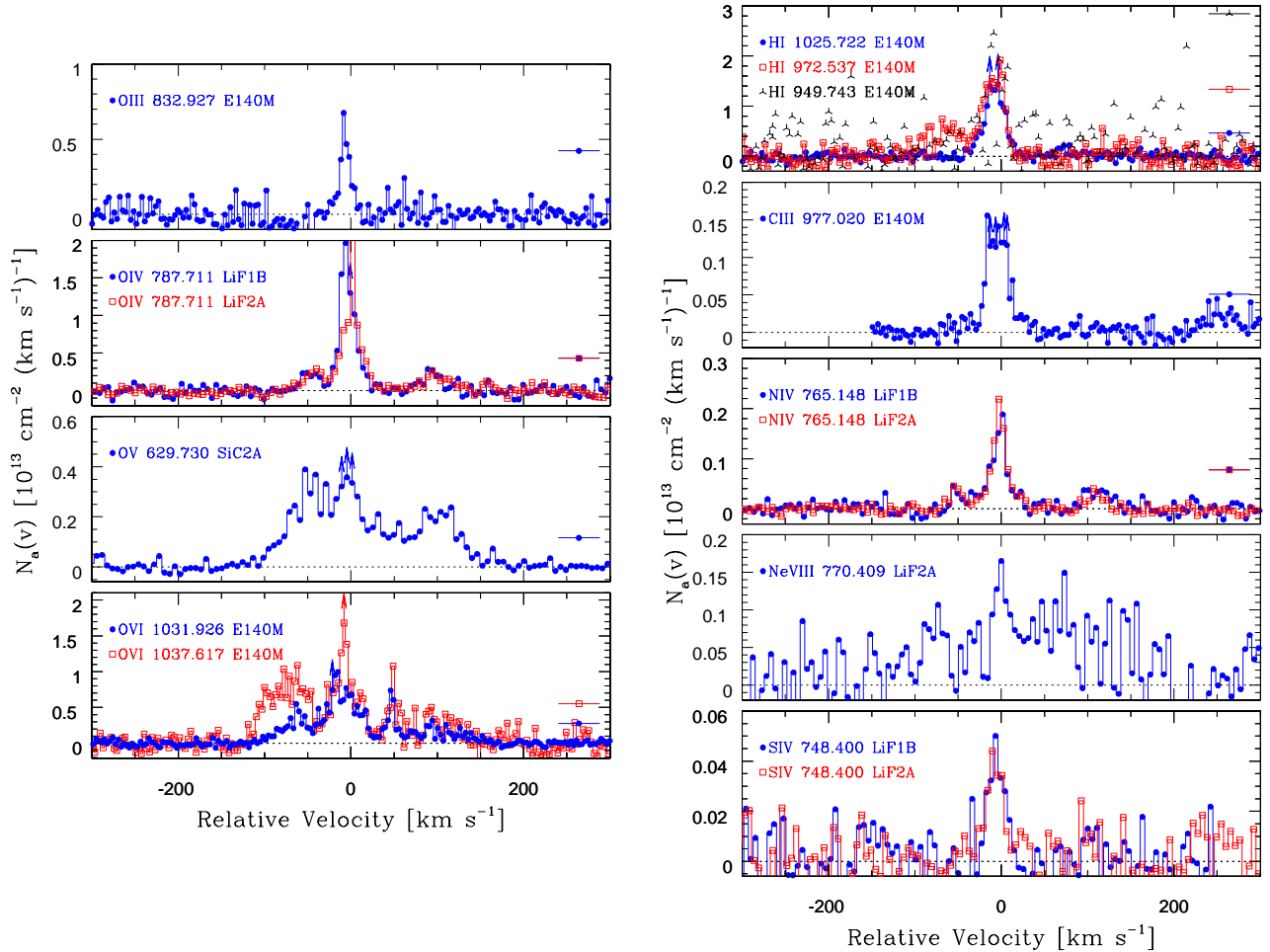


FIG. 6.— Apparent column density profiles of the detected species as a function of velocity relative to $z_{\text{abs}} = 0.4925$. Transitions from a given species are overplotted in the same panel. A horizontal bar on the right side of each panel marks the apparent column density where each transition reaches unit optical depth. FUSE data are shown with a sampling of 3 bins per resolution element, while STIS data are shown with a sampling of 2 bins per resolution element. Species of oxygen are plotted on the left, while other species are plotted on the right.

4. ASSOCIATED ABSORBER PROPERTIES AT $z_{\text{abs}} = 0.4925$

For associated absorption, our *HST* STIS and *FUSE* spectra the spectra cover the rest-frame wavelength range 610–1140 Å. In this range, we detect absorption at $z_{\text{abs}} = 0.4925$ in H I Lyman β , γ , δ , C III $\lambda 977.020$, N IV $\lambda 765.148$, O III $\lambda 832.927$, O IV $\lambda 787.711$, O V $\lambda 629.730$, O VI $\lambda \lambda 1031.926$, 1037.617 , Ne VIII $\lambda 770.409$, and S IV $\lambda 748.400$. Other transitions, notably C II $\lambda 1036.337$, O I $\lambda 877.879$, and O II $\lambda \lambda 834.466$, 833.329 are also covered in our spectrum, but are not detected. O III $\lambda 702.332$ is also covered by our spectrum and has a line strength which is larger than the 832.927 Å transition, but it is blended with Galactic Ar I $\lambda 1048.220$. [For this study, we adopt transition wavelengths and oscillator strengths from Morton (2003) for transitions above 912 Å, and from Verner et al. (1994) for transitions below 912 Å.] In Figure 5, we plot the detected transitions (histogram) in velocity with $z_{\text{abs}} = 0.4925$ defining the velocity zero-point. (Integration of the O III $\lambda 832.927$ line yields a centroid redshift of $z_{\text{abs}} = 0.49246 \pm 0.00006$. Henceforth, all quoted velocities are referenced to $z_{\text{abs}} = 0.4925$.) The

smooth curves overplotted on the data indicate our estimated continua obtained by fitting a low-order (≤ 5) Legendre polynomial to the adjoining continuum regions. These continuum fits follow the formalism described by Sembach & Savage (1992).

We make two remarks regarding the continuum fits to the O III $\lambda 832.927$ and Ne VIII $\lambda 770.409$ lines. Inspection of the continuum fit to the O III $\lambda 832.927$ lines shows several bins blueward of the absorption profile (corresponding to the velocity range $-100 \leq v[\text{km s}^{-1}] \leq -50$) that appear above the continuum. In order to fit these bins and lower them to the continuum level, we would require a polynomial order of at least 7, larger than what is prescribed by Sembach & Savage (1992). Using such a high order, we would be in danger of fitting out some of the absorption. Outside of the velocity range $-100 \leq v[\text{km s}^{-1}] \leq +5$ the fit appears quite good. In addition, the fit to the broad and shallow trough of the Ne VIII $\lambda 770.409$ is not well-represented in the expand view of Figure 5. We refer the reader to the lower left panel of Figure 2 where we show a larger wavelength range with the Legendre-polynomial fit to the local continuum superimposed. The continuum placement is rea-

TABLE 2. INTEGRATION OVER NARROW COMPONENTS (IN THE $z_{\text{abs}} = 0.4925$ REST-FRAME)

Ion	Wavelength ^a (Å)	$\log f\lambda^a$	$-30 \leq v[\text{km s}^{-1}] \leq +5$		$-30 \leq v[\text{km s}^{-1}] \leq +30$		$+30 \leq v[\text{km s}^{-1}] \leq +60$	
			W_λ (mÅ)	$N_a^{b,c}$ (10^{13} cm^{-2})	W_λ (mÅ)	N_a (10^{13} cm^{-2})	W_λ (mÅ)	N_a^d (10^{13} cm^{-2})
H I	1025.722	1.909	149 ± 4	≈ 42	< 21	< 1.8
	972.537	1.450	101 ± 5
	949.743	1.122	48 ± 8
C III	977.020	2.869	125 ± 7	≈ 6.2	< 26	< 0.26
N IV	765.148	2.684	...	$\lesssim 2.1$	119 ± 4	4.6 ± 0.3	< 12	< 0.26
O I	877.879	1.714	< 19	< 2.8	< 17	< 2.8
O II	833.329	2.097	< 20	< 1.3	< 19	< 1.2
O III	832.927	1.950	54 ± 5	8.4 ± 1.1	< 18	< 1.9
O IV	787.711	1.942	...	$\lesssim 11$	142 ± 5	36 ± 3	< 15	< 1.8
O V	629.730	2.511	164 ± 4	15 ± 2	66 ± 4	$\lesssim 2.9$
O VI	1031.926	2.136	238 ± 6	31 ± 2	88 ± 5	8.7 ± 1.2
Ne VIII	770.409	1.908	36 ± 6	5 ± 0.8	17 ± 4	$\lesssim 1.8$
S IV	748.400	2.573	...	$\lesssim 0.56$	28 ± 6	0.9 ± 0.2	< 14	< 0.35

 NOTE. — Errors on all quantities are quoted at 1σ confidence. Upper limits are quoted at 3σ confidence. The velocity range is quoted in km s^{-1} relative to $z_{\text{abs}} = 0.4925$.

^aAtomic data for transitions above the Lyman limit at 912 Å were taken from Morton (2003). For transitions blueward of the Lyman limit, we use data compiled by Verner et al. (1994).

^bThe column densities for H I and C III were determined using the integrated column density from the O III $\lambda 832.927$ profile and the column density ratios in Figure 7.

^cThe column density limits for N IV, O IV, and S IV were determined using the integrated column density from the O III $\lambda 832.927$ profile and the column density limits in Figure 7.

^dThe column density limits for O V and Ne VIII were determined using the integrated column density from the O VI 1037.617 Å profiles and the limiting column density ratios in Figure 10.

sonable.

In order to remove biases due to transition strengths ($\log f\lambda$) and compare profiles of species on a common (and physically meaningful) scale, we transform the flux profiles of each detected transition into apparent column density profiles using the method described by Savage & Sembach (1991).

$$N_a(v) = \frac{m_e c}{\pi e^2 f \lambda_o} \ln \left[\frac{I_c(v)}{I(v)} \right], \quad (2)$$

where f and λ are the oscillator strength and rest wavelength of the transition, and $I_c(v)$ and $I(v)$ are the continuum and attenuated fluxes, respectively. These are shown in Figure 6. Apparent column density profiles from multiple transitions from a given species are useful in testing for unresolved saturated structure, blends with transitions of other species from other absorption-line systems, and for assessing the levels of unocculted flux (e.g., scattered light within the instrument, or elevated background levels). For systems that are potentially intrinsic to the quasar, apparent column density profiles provide a means of assessing the fraction(s) of the quasar continuum/emission line region that is occulted by the absorbing gas (e.g., Scott et al. 2004b; Ganguly et al. 1999; Hamann et al. 1997; Barlow & Sargent 1997).

For the associated absorption along this sight line, our spectra cover multiple transitions of H I (Lyman β and higher order Lyman series), and O VI (the resonant doublet $\lambda\lambda 1031.927, 1037.617$). In Figure 6, we overlay the apparent column density profiles of the O VI doublet (left side, bottom panel). The apparent excess of column density in the O VI $\lambda 1037.617$ transition over the relative velocity range $-120 \leq v[\text{km s}^{-1}] \leq 0$ is due to

Galactic and high velocity cloud absorption by the C IV $\lambda 1548.204$ line (Fox et al. 2005). At all other velocities, the O VI $\lambda 1031.927$ and $\lambda 1037.617$ lines are in excellent agreement (within the errors) and there is no evidence for unresolved saturated structure, or for partial coverage of the quasar continuum/broad emission line region. Also, strongly saturated lines such as C III $\lambda 977.020$ are black in the line core (i.e., no flux is significantly detected), which further indicates that the absorber fully covers the continuum flux source.

There is a weak intervening IGM Ly α line at $z_{\text{abs}} = 0.19374$ (Lehner et al. 2006) that is blended with the Ly γ line from the associated absorber at $v \sim -65 \text{ km s}^{-1}$. The intervening Ly α line appears somewhat broad and shallow, and the Ly γ line of the associated absorber lies on its wing. Due to the noisiness of the data on the blue wing of the $z=0.19374$ Ly α line, it is difficult to assess its contribution to the $z=0.4925$ Ly γ profile, and the level of saturation in the core of the Lyman β . (The Ly δ line is detected, but also too noisy for this purpose.) We return to this issue in §4.1 by using the O III $\lambda 832.926$ profile. We note here that Lehner et al. (2006) report $\log N(\text{H I})[\text{cm}^{-2}] = 13.20 \pm 0.06$ and $b(\text{H I}) = 28.7 \pm 6.0 \text{ km s}^{-1}$ for the intervening $z_{\text{abs}} = 0.19374$ Ly α line.

We make one final note regarding the apparent column density profiles as a whole before continuing with more detailed investigations of the various components. The rest-frame transitions in the $732\text{--}792 \text{ Å}$ wavelength range lie in both the LiF1B and LiF2A detector segments of the *FUSE* spectra. We overlay the apparent column density profiles from both channels for the S IV $\lambda 748.400$ (Figure 6, right side, bottom panel), N IV $\lambda 765.148$ (Figure 6,

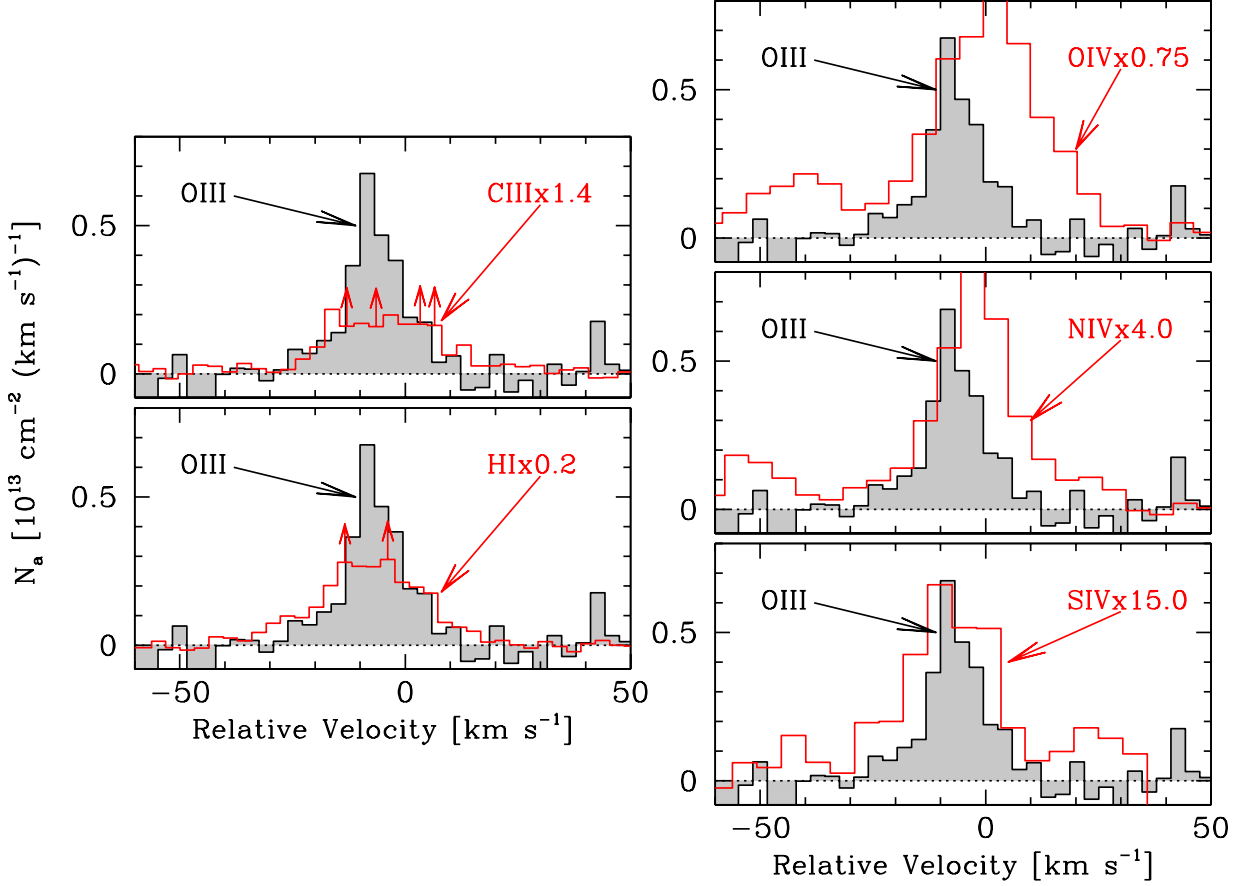


FIG. 7.— In the above panels, we overlay the O III apparent column density profiles (shaded histogram) on top of the H I, C III, N IV, O IV, and S IV apparent column density profile. Vertical arrows are shown for pixels where the flux is negative due to Poisson fluctuations at low flux levels. The H I and C III profiles are scaled to match the O III profile in unsaturated regions. The N IV and O IV profiles are scaled to match the blue wing of the O III profile. The S IV profile is scaled to match the peak of the O III profile.

right side, middle panel), and O IV $\lambda 787.711$ (Figure 6, left side, second panel from the top) lines. The agreement between the LiF1B and LiF2A profiles for N IV and S IV is excellent. For the O IV profiles, there is a discrepancy in the bins near 0 km s^{-1} where the line profile reaches zero flux. The profiles are in excellent agreement otherwise, and we have no reason to suspect either one. We attribute the discrepancy near 0 km s^{-1} to Poisson errors, and use the profiles from the LiF2A detector segment for our investigation below (due the higher signal-to-noise). The excellent agreement evident in Figure 6 indicates that the LiF1 and LiF2 channels have very similar spectral resolutions.

4.1. Narrow Components

4.1.1. Low and moderate ionization gas at $v = -8 \text{ km s}^{-1}$ (in the $z = 0.4925$ rest-frame)

In Figures 5–6, absorption is clearly detected at -8 km s^{-1} in a wide variety of ionization species, from neutral H I, to low-ionization C III and O III, to moderate-ionization N IV, O IV, S IV, to high-ionization O V, O VI and Ne VIII. The kinematics of the absorption at this velocity is complex (and possibly multi-phase). Neutral (H I) and low-ionization species (C III, O III) appear as a discrete feature over the velocity

range $-30 \leq v [\text{km s}^{-1}] \leq +5$, while the higher-ionization species (N IV, O IV) appear as a broader feature over the velocity range $-30 \leq v [\text{km s}^{-1}] \leq +30$. The more highly-ionized species (O V–VI, Ne VIII) appear absorbed with little substructure over the entire velocity range $-30 \leq v [\text{km s}^{-1}] \leq +30$, not as a single discrete feature. We focus here on the discrete features appearing in the H I, C III, N IV, O III–IV, and S IV ions.

To investigate the number of phases that are required to explain the absorption at -8 km s^{-1} and the ionization conditions of each phase, we first overlay the apparent column density profiles of the neutral, low- and moderate-ionization species in Figure 7. In the left panels of Figure 7, we overplot the H I and C III apparent column density profiles on top of the O III profile. For H I, we use the Ly β profile because it provides the best and cleanest information in the shapes of the wings without suffering from blends (as in the case of Ly γ) or noise (as in the case of Ly δ). We comment on this further below. The kinematics of these species is remarkably similar and suggest an origin in the same phase. (Note the gradual trend toward increasing column density in the velocity range $-30 \leq v \leq -15 \text{ km s}^{-1}$, and the sharp drop-off at 5 km s^{-1} .) The reason for this asymmetry is unclear. If the absorption results from

the superposition of symmetric (e.g., Gaussian) profiles, then at least two such components are needed to reproduce the observed profile. However, it is not clear that such an assumption is required (e.g., Ganguly et al. 2005; Arav, Korista, & de Kool 2002). Regardless, if the kinematics of the H I and C III in the velocity range $-15 \leq v \leq +5 \text{ km s}^{-1}$ are also well-traced by the O III profile, we can use the O III profile as a template to estimate the level of saturation, and, more importantly, to obtain column density ratios to constrain the ionization conditions of the low-ionization phase. We summarize these ratios below.

In the right panels of Figure 7, we overplot the O IV, N IV, and S IV profiles on top the O III profile. These overlays also show that the kinematics of these profiles do not match that of the O III profile. They are wider, possibly resulting from a blend of different kinematic components (see §4.1.2. With the simple assumption that any O IV, N IV and S IV that co-exists in the O III-bearing gas have the same kinematic profile, we can place limits on how much of these higher-ionization species relative to O III could arise in that phase (in the interest of providing the most information to constrain ionization models). (Although there is no apparent shift in the centroid of S IV $\lambda 784.400$ line relative to the O III $\lambda 832.927$ line, the kinematics of the two profiles do not match. While it is possible that this could be an effect of the resolution, since the S IV $\lambda 784.400$ line appears in the FUSE LiF2 channel, the profile spans over three resolution elements. Thus, we report only a limit on the $N(\text{S IV})/N(\text{O III})$ ratio.) In Figure 7, we use the O III profile, which appears over the velocity range $-30 \leq v[\text{km s}^{-1}] \leq +5$, to estimate the following constraints on the column density ratios for the O III-bearing gas:

$$\begin{aligned} N(\text{H I})/N(\text{O III}) &\approx 5 \\ N(\text{C III})/N(\text{O III}) &\approx 0.74 \\ N(\text{N IV})/N(\text{O III}) &\lesssim 0.25 \\ N(\text{O IV})/N(\text{O III}) &\lesssim 1.3 \\ N(\text{S IV})/N(\text{O III}) &\lesssim 0.07 \end{aligned} \quad (3)$$

In columns 4 and 5 of Table 2, we report our equivalent width measurements and further summarize our column density assessments of the O III-bearing gas using the above ratios. Direct integration of the O III profile over the above velocity range yields a column density of $N_{\text{a}}(\text{O III}) = (8.4 \pm 1.1^{+0.5}_{-1.7}) \times 10^{13} \text{ cm}^{-2}$, where the first error is the 1σ confidence error resulting from statistical and continuum placement uncertainties and the latter error is the systematic uncertainty resulting from the choice of integration range. (To assess this latter error, we added and subtracted 10 km s^{-1} to the integration range.) We note that the resulting H I column density implied by eq. 3 [$N(\text{H I}) \approx (4.2 \pm 0.6^{+0.3}_{-0.9}) \times 10^{14} \text{ cm}^{-2}$] is consistent with the curve of growth of the Lyman series equivalent widths. In addition, we integrated the region of the spectrum where the O II $\lambda 833.329$ line ($\log f\lambda = 2.097$) is expected and find a limiting column density of $N(\text{O II}) < 1.3 \times 10^{13} \text{ cm}^{-2} (3\sigma)$ for this low-ionization phase.

To understand the ionization structure of the O III-bearing phase, we assume that the absorption arises from a photoionized plane-parallel slab with a solar relative abundance pattern [$n(\text{C})/n(\text{O}) = 0.537$;

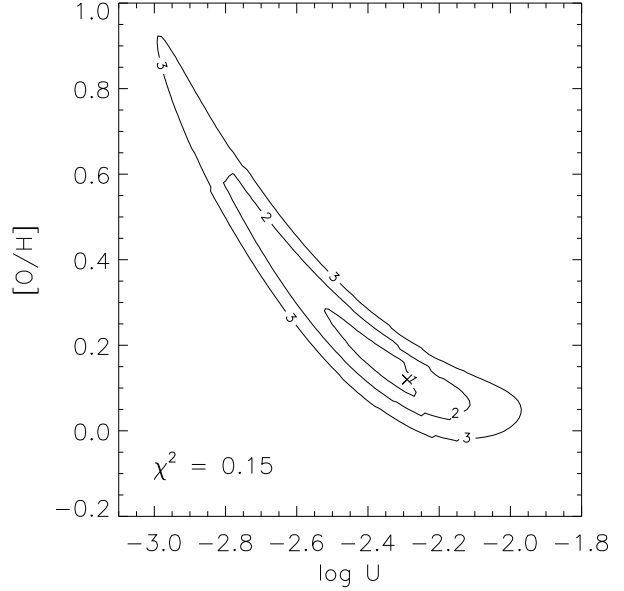


FIG. 8.— In the plot we show $\Delta\chi^2$ contours on the metallicity ($[\text{O}/\text{H}]$) - ionization parameter (U) plane for the low-ionization phase of the $v = -8 \text{ km s}^{-1}$ component probing the narrow emission line region. The optimal parameters are marked with a cross while contours are drawn for 1, 2, and 3σ confidence regions.

Asplund, Grevesse, & Sauval (2005), and references therein] and use the above column density ratios to infer the ionization parameter, U , defined as the number density of ionizing photons per baryon, and metallicity, $[\text{O}/\text{H}]$. We use the CLOUDY photoionization code (Ferland et al. 1998) to compute the best-fit model, specifying the radiation field as an AGN spectrum with the following parameters (as described in §3): $T = 15,000 \text{ K}$, $\alpha_{\text{ox}} = -1.46$, $\alpha_{\text{UV}} = -0.58$, $\alpha_{\text{x}} = -2.1$, normalized to $\log L_{\nu}(1 \text{ Ryd}) = 30.65$. (See §3 for a justification of these parameters.) In Figure 8, we show χ^2 contours on the $[\text{O}/\text{H}] - U$ plane. The best fitting parameters are:

$$\begin{aligned} [\text{O}/\text{H}] &= +0.12^{+0.16}_{-0.03}, \\ \log U &= -2.29^{+0.02}_{-0.23}, \\ \log N(\text{H}) &= 17.54^{+0.04}_{-0.25}. \end{aligned} \quad (4)$$

The total hydrogen column density of the O III-bearing phase results from a combination of the integrated O III column density (Table 2), the H I/O III column density ratio from equation 3, and the H I ion fraction implied by the ionization parameter: $N(\text{H}) = N(\text{O III}) \times [N(\text{H I})/N(\text{O III})] \times f_{\text{H I}}^{-1}(U)$. The ionization parameter is most tightly constrained by the O IV/O III limit and C III/O III (with the assumed solar relative C/O abundance) ratio.

Increasing the carbon to oxygen relative abundance to $[\text{C}/\text{O}] = +0.1$ yields a slightly better fit to the observations (best parameters: $[\text{O}/\text{H}] \sim +0.45$, $\log U \sim -2.73$), since the lower constraint on the ionization parameter provided by C III/O III is relaxed. The best fit model presented above yields O IV/O III at the observational limit, while the $[\text{C}/\text{O}] = +0.1$ best-fit model predicts O IV/O III well below the limit. (The CLOUDY model produces a factor of 1.3 less N IV column density than the observational limit.) All other best-fit models with

non-solar carbon to oxygen abundances ($[C/O] > 0.1$, $[C/O] < 0$) fail to reconcile the O III column density and C III/O III column density ratio. (Photoionization models with $[C/O] > 0.1$ predict ratios that are too large; those with $[C/O] < 0$ predict ratios that are too small.) Such an enhancement may be reasonable given that the Fe/ α relative abundance appears to be enhanced in the emission-line regions of AGN (e.g., Hamann & Ferland 1999, and references therein). There is evidence from the atmospheres of Galactic halo stars (e.g., McWilliam 1997) that the C/Fe is roughly constant (at the solar value) over a large range in metallicity. If this insensitivity can be extrapolated to the super-solar metallicities of AGN line-emitting regions, then a slightly super-solar C/O relative abundance in this $z_{\text{abs}} \sim 0.5$ absorber is reasonable.

We note here that the ionization fractions of O V and O VI for the O III-bearing gas are 6.75×10^{-2} and 2.52×10^{-3} , respectively, compared to 0.4 for O III. This implies column densities of $N(\text{O V}) \sim 10^{13} \text{ cm}^{-2}$, and $N(\text{O VI}) \sim 5 \times 10^{11} \text{ cm}^{-2}$. Integration of the O V and O VI over the velocity range $-30 \leq v [\text{km s}^{-1}] \leq +5$ yield apparent column densities of $N_a(\text{O V}) = (1.1 \pm 0.1) \times 10^{14} \text{ cm}^{-2}$, and $N_a(\text{O VI}) = (2.4 \pm 0.4) \times 10^{14} \text{ cm}^{-2}$. These high-ionization species must arise in a phase/component of gas that is separate from the O III-bearing gas.

4.1.2. Moderate and High ionization gas at $v = 0 \text{ km s}^{-1}$ (in the $z = 0.4925$ rest-frame)

As we pointed out in the previous section from the overlays of the O IV, N IV, and S IV apparent column density profiles on top of the O III profile, we can only place limits on the amount of higher-ionization species that co-exists in the O III-bearing gas. We return to the issue of the higher-ionization species here. These higher-ionization species appear over the velocity range $-30 \leq v [\text{km s}^{-1}] \leq +30$. In columns 6 and 7 of Table 2 we report equivalent width and apparent column density integrations of these and high-ionization species (O V-VI, Ne VIII) over this velocity range. (We do not report additional measurements of H I, C III, or O III as there is no additional information over that reported in columns 4 and 5.)

As we pointed out earlier, there are differences in the apparent velocity widths of the O IV, N IV, and S IV compared to O III, H I, and C III. There are two additional aspects of Figure 7 that merit consideration. First, there is a $\sim +8 \text{ km s}^{-1}$ velocity shift in the peak of the O IV and N IV profiles relative to the O III profile. Second, in the velocity range $+5 \leq v \leq +30 \text{ km s}^{-1}$, there is significant column density apparent in O IV, whereas none exists in the O III, H I, and C III profiles.

From this, we conclude that there is at least one additional higher-ionization phase of gas at these velocities which produces O IV, N IV, and possibly some O V, and O VI. This “component” is not necessarily restricted to the velocity range $+5 \leq v [\text{km s}^{-1}] \leq +30$, however. It is certainly plausible that the absorption by this gas extends down to velocities where the O III-bearing gas appears. (For this reason, we can only quote limits on the column density ratios in the previous section.) In fact, as mentioned in the previous section, additional gas in the velocity range common with that of the O III-bearing gas (i.e., $-30 \leq v [\text{km s}^{-1}] \leq +5$) is required to explain the

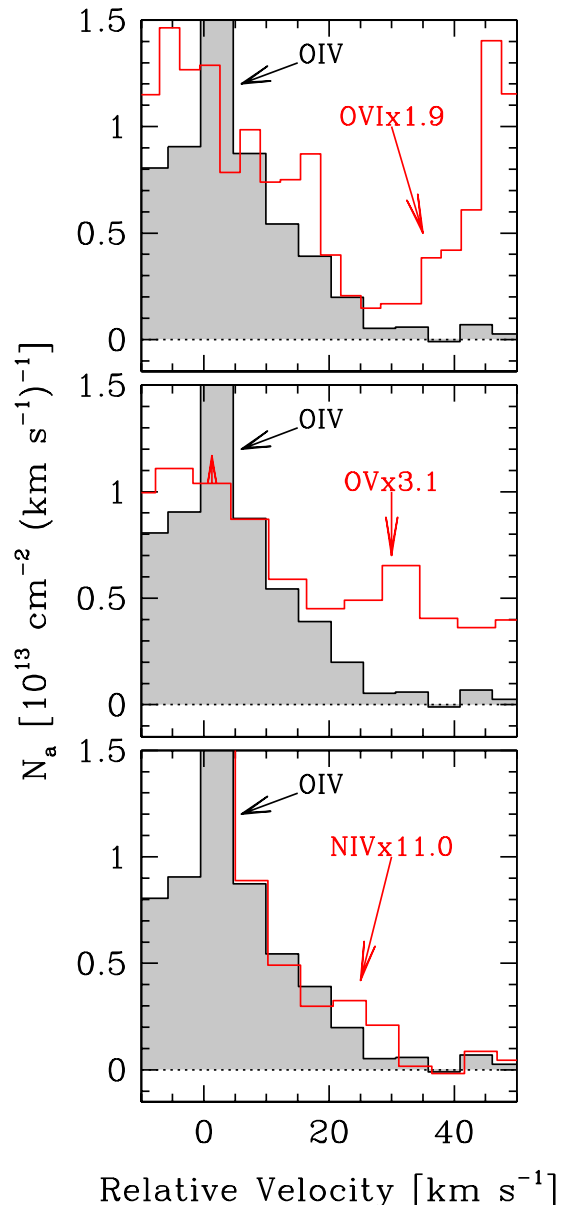


FIG. 9.— In the figure, we overlay the apparent column density profile of O IV (shaded histogram) on top of the O VI (top panel), O V (middle panel), and N IV (bottom panel) profiles (unshaded histograms) in the region of the $v = 0 \text{ km s}^{-1}$ component. The N IV profile has been scaled to match the O IV profile in the velocity range $+5 \leq v [\text{km s}^{-1}] \leq +30$. The O V-VI profiles have been scaled to the minimum allowable value to hide the O IV profile in that same velocity range.

column densities of the high-ionization species O V-VI.

Since much of the O IV and N IV column density in the velocity range $-30 \leq v [\text{km s}^{-1}] \leq +5$ can be explained by the O III-bearing gas, we focus here on the remaining velocity range $+5 \leq v [\text{km s}^{-1}] \leq +30$ (with the aforementioned caveats). In Figure 9, we compare the apparent column density profiles of O V-VI and N IV to that of O IV. In the velocity range $+5 \leq v [\text{km s}^{-1}] \leq +30$, there is good agreement between the kinematics of O IV and N IV, so the two likely co-exist in the same phase of gas. [Note that the N IV profile has been scaled vertically.] Figure 9 also shows that the kinematics of O V-VI do not

match O IV in the velocity range $+5 \leq v[\text{km s}^{-1}] \leq +30$. This implies that the absorbing gas of those two species at these velocities arise in a separate phase/component (probably one of the broad components described in §4.2). Using the comparison of the apparent column density profiles we can still place limits on how much O V-VI could arise in the O IV-bearing gas at these velocities. [Here, we again make the simple assumption that any O V-VI that co-exists in the O IV-bearing gas must have the same kinematics as the O IV over those velocities.] We find the following constraints:

$$\begin{aligned} N(\text{N IV})/N(\text{O IV}) &\approx 0.1 \\ N(\text{O V})/N(\text{O IV}) &\lesssim 0.32 \\ N(\text{O VI})/N(\text{O IV}) &\lesssim 0.53 \end{aligned} \quad (5)$$

For the O IV-bearing gas in the velocity range $+5 \leq v[\text{km s}^{-1}] \leq +30$, we find $N_a(\text{N IV})/N_a(\text{O IV}) \approx 0.1$ (compared to ≈ 0.2 for the CLOUDY model of the O III-bearing gas). In the photoionization models described above, the smallest achievable $N(\text{N IV})/N(\text{O IV})$ ratio is ≈ 0.14 , which occurs at an ionization parameter of $\log U \approx -2.6$. This is problematic for two reasons. First, at this ionization level, we would expect to see O III at these velocities which we do not (see Figures 6 and 7) down to a 3σ confidence limit of $1.6 \times 10^{13} \text{ cm}^{-2}$. Second, while the ratio is close, this ionization level does not really explain the observed ratio. There are two possible conclusions from this simple analysis. Either pure photoionization by the quasar may be the only ionization mechanism for the O IV/N IV-bearing gas at these velocities, or the N/O relative abundance deviates from our assumed solar ratio ($\text{N/O} = 0.132$, Asplund et al. 2005).

If the ionization mechanism of this gas is predominantly photoionization by the quasar, then the lack of O III implies that the ionization parameter must be larger than $\log U \approx -1.7$. At this ionization level, and with solar relative abundances, the N IV/O IV ratio is ≈ 0.17 . Thus, the N/O abundance must be lowered by at least ≈ 0.07 dex relative to the solar value (i.e., $[\text{N/O}] \lesssim -0.07$) to match the observed N IV/O IV ratio.

4.1.3. High-ionization gas at $v = +50 \text{ km s}^{-1}$ (in the $z = 0.4925$ rest-frame)

In Figures 5-6, there is clearly a narrow component detected in O VI at $+50 \text{ km s}^{-1}$ in the $z = 0.4925$ rest-frame. This component is not detected in any low- to moderate-ionization species (i.e., H I, C III, N IV, O III, O IV, S IV). Absorption is detected at this velocity in O V and Ne VIII. However, it appears that those species arise from the broader component at $v \sim 100 \text{ km s}^{-1}$. If so, the only ion detected in this narrow component is O VI. To place constraints on the column densities of species (for the consideration of ionization models below), we integrate the apparent column density profiles of non-detected species and O VI over the velocity range $+30 \leq v \leq +60 \text{ km s}^{-1}$. For O V and Ne VIII, we take the same approach as in the previous section and consider how much of the absorption observed in this velocity range could be attributed to the O VI component. In Figure 10, we plot the O VI apparent column density profile over those of O V (bottom panel) and Ne VIII (top panel). In each case, we have scaled the apparent column density profile to match the level of the O VI.

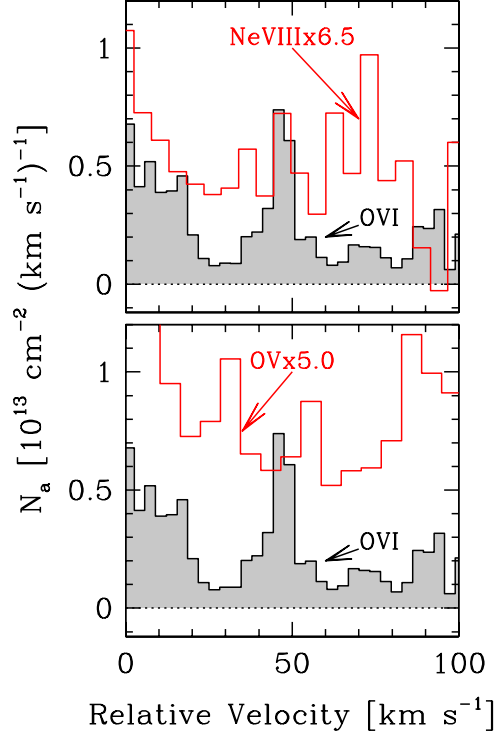


FIG. 10.— In the figure, we overlay the apparent column density profile of O VI (shaded histogram) on top of the O V (bottom panel) and Ne VIII (top panel) profiles (unshaded histograms) in the region of the $v = +50 \text{ km s}^{-1}$ component. The O V and Ne VIII profiles have been scaled to the minimum allowable value to hide the O VI profile.

This yields the following constraints:

$$\begin{aligned} N(\text{O V})/N(\text{O VI}) &\lesssim 0.2 \\ N(\text{Ne VIII})/N(\text{O VI}) &\lesssim 0.15 \end{aligned} \quad (6)$$

In columns 8 and 9 of Table 2, we summarize our equivalent width measurements (over the velocity range $+30 \leq v[\text{km s}^{-1}] \leq +60$) and column density constraints.

The velocity width of O VI ($b = 9 \pm 1 \text{ km s}^{-1}$)⁶ implies a maximum temperature of $\log T(\text{K}) \leq 4.9$. At this temperature collisional processes do not produce appreciable amounts of O VI, so this high-ionization gas likely arises from photoionization. The central engine is likely to dominate the radiation field so we use the same ionizing spectrum and range of ionization parameters used in the previous two sections. The location of the gas is not clear since there is no additional information (such as a velocity coincidence with a line-emitting region, or a density diagnostic). As a result there is fundamental uncertainty in the metallicity of the gas and, by consequence, the total hydrogen column density. [Since H I is not detected down to $N(\text{H I}) < 1.8 \times 10^{13} \text{ cm}^{-2}$ (see Table 2), we do not have a means of inferring the metallicity. We can constrain the metallicity only with a model for the ionization, and we return to this below.] To place ionization constraints on the gas, we ran a grid

⁶ This b -value was obtained by taking the second moment of the apparent optical depth distribution over the velocity range $+30 \leq v[\text{km s}^{-1}] \leq +60$.

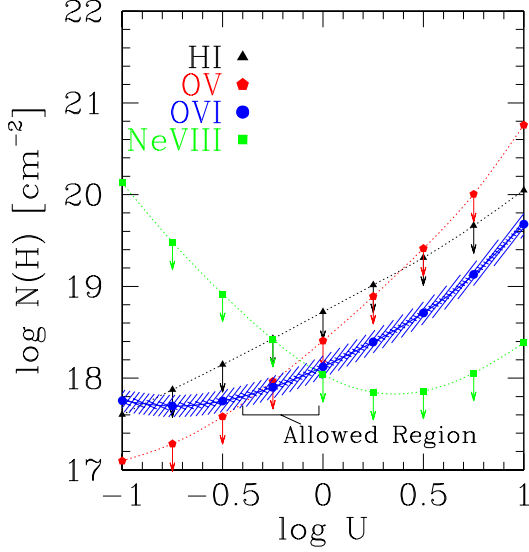


FIG. 11.— The figure shows isopleths of species column density in total column density $[N(\text{H})]$ -ionization parameter (U) space for the $v = +50 \text{ km s}^{-1}$ component. The relevant species for each curve is distinguished by the symbols shown in the corner. For O VI, the hatched regions around the curve indicate the 1σ and 3σ confidence uncertainties. For all other species, the curves are shown as upper limits. The region of parameter space that satisfies all column density information is marked “Allowed region.” The models assumed a metallicity $[\text{O}/\text{H}] = +0.12$, with solar relative abundances.

of CLOUDY models over a range of ionization parameters ($-1 \leq \log U \leq 1$) using the same ionizing spectrum and metallicity ($[\text{O}/\text{H}] = 0.12$) as in the previous section for the O III phase. For each species, we generated a curve of total hydrogen column density (as a function of ionization parameter) which reproduces the observed species column density (or limit). The most constraining curves (those for H I, O V, O VI, and Ne VIII) are shown in Figure 11. The shapes of the curves are determined by the ionization fraction of the species, with the minimum in each curve occurring at the ionization parameter that maximizes that species. The relative vertical placements of the curves are determined by the relative abundances of the elements (which are assumed to be solar in these models) except for H I. If the gas can be described by single photoionized phase, then there must be a location (or region) on the $N(\text{H})$ - U plane that satisfies all column density measurements and constraints. In the case of this component, such a region does exist:

$$\begin{aligned} (17.65)17.75 \lesssim \log N(\text{H}) \lesssim 18.1(18.15) \\ (-0.45) - 0.35 \lesssim \log U \lesssim -0.02(0). \end{aligned} \quad (7)$$

The quoted range reflects the 1σ confidence uncertainty resulting from the error in $N(\text{O VI})$. The parenthetical numbers indicate the 3σ confidence range. The lower limit on the ionization parameter is established by the comparison of $N(\text{O VI})$ and the limit on $N(\text{O V})$. This is independent of assumed metallicity or relative abundance. The upper limit arises from comparing the O VI and Ne VIII column densities which relies on the assumed relative Ne/O abundance (taken to be $\text{Ne}/\text{O} = 0.151$, Asplund et al. 2004, and references therein). The limits on the ionization parameter are probably not sensitive to the uncertainty in the metallicity since the absorber is

optically thin (in spite of the large total hydrogen column density) to the ionizing radiation.

If the derived range of ionization parameter for the gas is accurate, then we can place a lower limit to the metallicity of the gas given the O VI column density and the H I column density non-detection limit via:

$$\begin{aligned} [\text{O}/\text{H}] &= \log \left[\frac{N(\text{O})}{N(\text{H})} \right] - \log \left[\frac{N(\text{O})}{N(\text{H})} \right]_{\odot} \\ &= \log \left[\left(\frac{N(\text{O VI})}{N(\text{H I})} \right) \left(\frac{f_{\text{O VI}}^{-1}(U)}{f_{\text{H I}}^{-1}(U)} \right) \right] + 3.34, \end{aligned} \quad (8)$$

where $f(U)$ is the ionization fraction and we have used the Asplund et al. (2004) value for the solar oxygen abundance. The minimum value for $f_{\text{O VI}}^{-1}(U)/f_{\text{H I}}^{-1}(U)$ arises at the largest allowed ionization parameter. For $\log U \sim -0.02$, this fraction is 2.2×10^{-5} . Using the integrated O VI column density and H I column density limit from Table 2, we find a limiting metallicity of $[\text{O}/\text{H}] \geq -0.62$.

For a given absorber density, the $N(\text{H})$ axis can also be interpreted as the thickness of the slab and U as the distance between the absorber and quasar central engine. Using the limiting metallicity, the total hydrogen column density could be as large as $N(\text{H}) \sim 5 \times 10^{18} \text{ cm}^{-2}$.

It is noteworthy that over the range of ionization parameters allowed by the column density constraints, the dominant stage of oxygen is O VII. For this component, our models predict O VII column densities in the range $\log N(\text{O VII}) = 14.4 - 14.8$.

4.2. Broad Components

On the blue and red wings of the associated absorption-line system, we detect broad components in the higher-ionization species: N IV, O IV-VI, and Ne VIII. Neither S IV, nor any species with an ionization stage smaller than four is detected. In Figure 12, we show overlays of the O IV-VI, N IV, and Ne VIII apparent column density profiles in the velocity ranges $-140 \leq v \leq -20 \text{ km s}^{-1}$ and $+60 \leq v \leq +200 \text{ km s}^{-1}$. In Table 3, we present measurements (equivalent widths and integrated apparent column densities) of the ions covered by our data over those two velocity ranges.

Obtaining velocity centroids and widths for these components in a non-parametric way using the apparent optical depth method is a difficult task for two reasons. Both of these quantities involve moments of the apparent optical depth and can be greatly affected by the noise when the absorption is relatively weak (e.g., with equivalent width less than five times the error), and by the choice of integration range (which itself is affected by the assumed structure of the absorption profile). For O V $\lambda 629.73$ and O VI $\lambda 1031.926$, which are the strongest unblended lines in these components, we find $b \sim 23 - 32 \text{ km s}^{-1}$ for the blue wing and $b \sim 29 - 40 \text{ km s}^{-1}$ for the red wing using the integration ranges quoted above and in Table 3. These ranges include an assessment of the uncertainty due to the choice of integration range by adding and removing 20 km s^{-1} to the range.

On the red wing, the apparent column density profile of O IV-VI appear to match, at least to within the errors. There is a possible excess of column density in the O V profile at $v = +110 \text{ km s}^{-1}$ over O VI, and an excess of

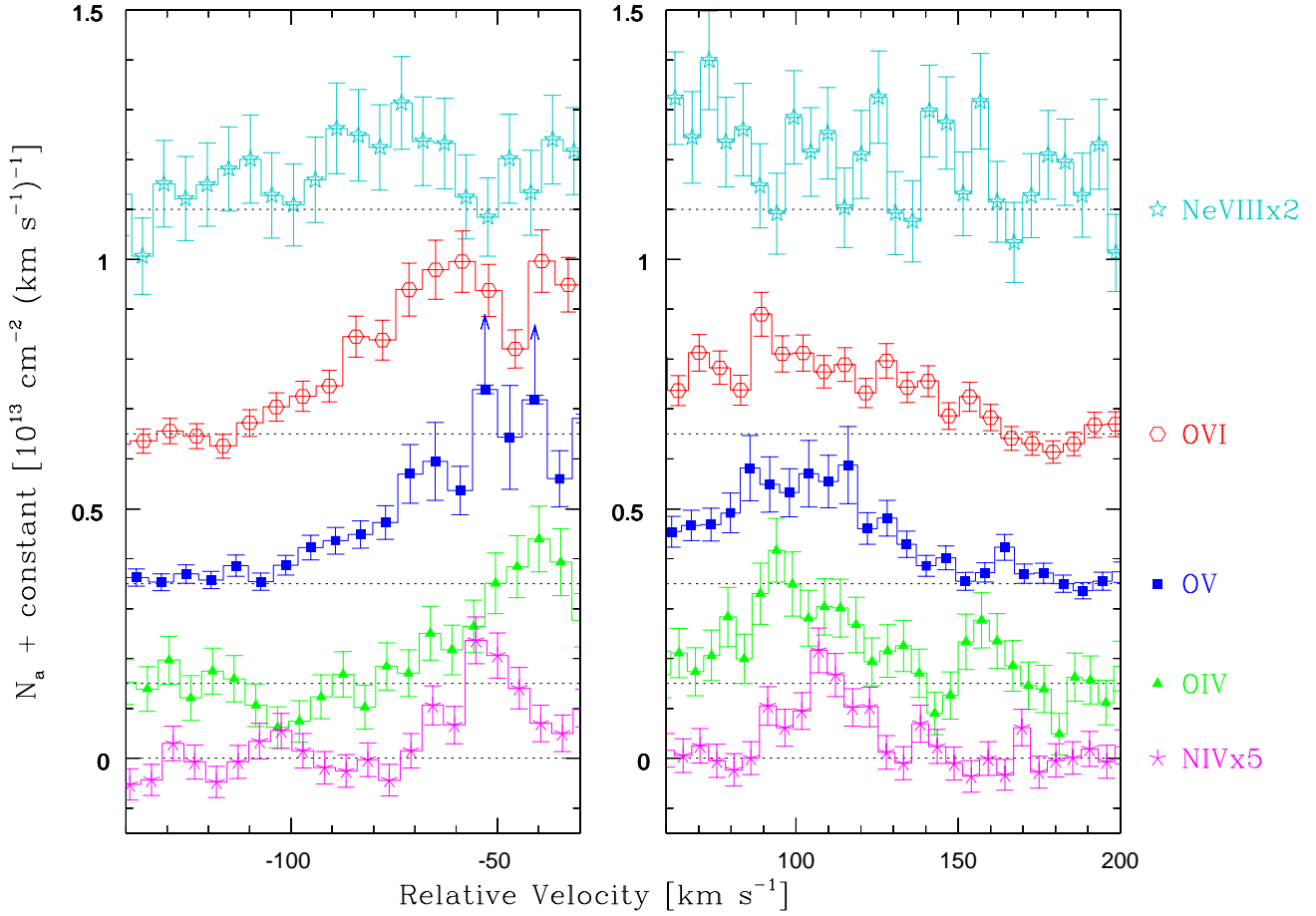


FIG. 12.— In the above panels, we overlay the velocity-aligned apparent column density profiles of the N IV, O IV-VI, and Ne VIII ions on blue and red wings (left and right panels, respectively) of the associated absorber. For purposes of clarity, the profiles are offset vertically, with the “zero-level” for each profile shown as a horizontal dotted line. The O VI profiles are shown with a sampling of 7 km s^{-1} per bin, similar to the sampling the FUSE data. In addition, the N IV and Ne VIII profiles are shown with a multiplicative factor to offset the abundance differences relative to oxygen.

TABLE 3. INTEGRATION OVER BROAD COMPONENTS (IN THE $z_{\text{abs}} = 0.4925$ REST-FRAME)

Ion	Transition ^a (Å)	$\log f\lambda^a$	$-140 \leq v [\text{km s}^{-1}] \leq -30$		$+60 \leq v [\text{km s}^{-1}] \leq +200$	
			W_λ (mÅ)	N_a (cm^{-2})	W_λ (mÅ)	N_a (cm^{-2})
H I	1025.722	1.909	< 44	$< 3.5 \times 10^{13}$	< 53	$< 4.3 \times 10^{13}$
C III	977.020	2.869	< 49	$< 4.5 \times 10^{12}$	< 69	$< 5.7 \times 10^{12}$
N IV	765.148	2.684	30 ± 8	$(8.0 \pm 1.7) \times 10^{12}$	40 ± 11	$(9.7 \pm 1.9) \times 10^{12}$
O I	877.879	1.714	< 36	$< 5.2 \times 10^{13}$	< 42	$< 5.9 \times 10^{13}$
O II	833.329	2.097	< 38	$< 2.5 \times 10^{13}$	< 46	$< 2.7 \times 10^{13}$
O III	832.927	1.950	< 42	$< 3.4 \times 10^{13}$	< 45	$< 3.9 \times 10^{13}$
O IV	787.711	1.942	40 ± 10	$(6.0 \pm 1.2) \times 10^{13}$	69 ± 11	$(9.2 \pm 1.4) \times 10^{13}$
O V	629.730	2.511	182 ± 10	$(1.5 \pm 0.2) \times 10^{14}$	206 ± 12	$(1.4 \pm 0.1) \times 10^{14}$
O VI	1031.926	2.136	201 ± 13	$(1.8 \pm 0.1) \times 10^{14}$	171 ± 16	$(1.2 \pm 0.1) \times 10^{14}$
Ne VIII	770.409	1.908	32 ± 8	$(4.2 \pm 1.0) \times 10^{13}$	51 ± 10	$(6.9 \pm 1.2) \times 10^{13}$
S IV	748.400	2.573	< 26	$< 6.7 \times 10^{12}$	< 31	$< 7.9 \times 10^{12}$

NOTE. — Errors on all quantities are quoted at 1σ confidence. Upper limits are quoted at 3σ confidence.

^aAtomic data for transitions above the Lyman limit at 912 Å were taken from Morton (2003). For transitions blueward of the Lyman limit, we use data compiled by Verner et al. (1994).

O VI at $v = +140 \text{ km s}^{-1}$ over O IV. The O IV profile is consistent with zero column density at $v = +140 \text{ km s}^{-1}$. It is not clear if this indicates an actual termination of the broad component (at a velocity smaller either the O V or O VI) with an additional component at higher velocities, or whether both are actually part of the same component. The similarity of the O V and the O VI profiles (modulo the aforementioned excess) indicates that both are produced in the same gas. It is difficult to place interesting ionization constraints on this phase given the lack of additional information.

On the blue wing, there are clear differences in the kinematics of the O IV profiles compared to O V and O VI, which appear to rule out the production of all three species in the same phase. The O IV profile appears to terminate at $v = -80 \text{ km s}^{-1}$, with no detectable column density blueward of that velocity. On the other hand, O V and O VI are detected at those velocities and appear to trace each other perfectly at $v \leq -90$, with the termination of the profile near $v = -110 \text{ km s}^{-1}$. In the range $-90 \leq v \leq -60 \text{ km s}^{-1}$, there is a clear systematic excess of O VI over O V. We note, however, that the O V $\lambda 629.730$ line becomes optically thick ($\tau = 1$) at an apparent column density of $1.2 \times 10^{12} \text{ cm}^{-2} (\text{km s}^{-1})^{-1}$, whereas the O VI $\lambda 1031.927$ line becomes optically thick at $2.8 \times 10^{12} \text{ cm}^{-2} (\text{km s}^{-1})^{-1}$. So the differences in the profiles may be due to the effects of unresolved saturation.

We assume from the kinematic similarity of the O V and O VI apparent column density profiles that the two species arise in the same gas. With only two ions with broad kinematics, however, it is difficult to motivate a single ionization mechanism. Thus, we consider both collisional ionization equilibrium (CIE) and photoionization equilibrium (PIE) processes. In either case, the ionization level (determined by temperature in CIE, or by ionization parameter in PIE) and total oxygen column density can be gauged from the O VI/O V column density ratio. Inspection of the integrated column densities in Table 3 reveals that the O VI/O V ratio is slightly higher on the blue wing than the red wing, implying slightly higher ionization, although in both cases the ratio is nearly unity. In CIE, this is easily achieved at a temperature of $\log T[\text{K}] \sim 5.5$. Photoionization by the quasar can also produce such a ratio with an ionization parameter of $\log U \sim -0.9$. We note that, in either case, thermal broadening produces a line width that is smaller than observed [$b_{\text{thermal}}(\text{O}) \lesssim 18 \text{ km s}^{-1}$ vs. $b \sim 23 - 40 \text{ km s}^{-1}$], so there are likely other contributions to the kinematics.

In Figure 13, we compare the predicted N IV, O IV, and Ne VIII apparent column density profiles from CIE and PIE models with the observed profiles. These profiles were generated by first assessing the temperature profile (for CIE) or ionization parameter profile (for PIE) that reproduces the O VI/O V ratio at each velocity bin. The apparent column density profiles were computed from the ionization fractions implied by the temperature/ionization parameter profiles and assumed solar relative abundances from Asplund et al. (2004, $\text{N/O}=0.132$, $\text{Ne/O}=0.151$):

$$N(X, i) = N(\text{O}, 5) \left(\frac{f(X, i)}{f(\text{O}, 5)} \right) \left(\frac{X}{\text{O}} \right), \quad (9)$$

where $N(X, i)$ is the column density (or apparent col-

umn density per unit velocity) of the i th ionization stage of element X, $f(X, i)$ is the ionization fraction of that ion, and (X/O) is the abundance of X relative to oxygen. (This is a more general recasting of equation 8.) All quantities are taken to be functions of velocity. As mentioned above, the ionization fractions are determined by temperature in CIE models, or by ionization parameter in PIE models. The temperature profile was computed using the collisional ionization equilibrium tables from Sutherland & Dopita (1993). The ionization parameter profile was computed using a grid of CLOUDY models assuming the quasar radiation field. In both ionization scenarios, detectable (and similar) amounts of O IV are produced, though not enough to fully explain the observed profile. Under the assumption of solar relative abundances, the CIE model does not produce N IV (at the requisite temperature to satisfy the O VI/O V ratio) while the PIE model produces small amounts. At least some of the O IV and N IV must reside in a separate phase of gas.

It is noteworthy that neither model of the O V-VI-bearing gas is able to produce any Ne VIII (under the assumption of solar relative abundances) as shown in the bottom panel of Figure 13. This may be an indication (in addition to the lack of kinematic similarities) that Ne VIII is produced in a completely separate, very highly ionized phase of gas. Such a phase would probably have a temperature of $\log T[\text{K}] \sim 5.85$, if collisionally ionized, or $\log U \sim +0.4$, if photoionized. At these ionization levels, all the oxygen would be in the form of O VII. Assuming the solar relative abundance, we predict column densities in the range $\log N(\text{O VII})[\text{cm}^{-2}] = 15 - 15.5$ for each of the two wings. If the metallicity of this very highly ionized phase is solar, then the total hydrogen column density for each wing is $\log N(\text{H})[\text{cm}^{-2}] = 18.3 - 18.8$. However, the origin of this gas, and therefore its metallicity, is unclear. It could arise near the quasar, or far from it in the halo of the host galaxy.

An alternative to a separate phase of gas that could potentially explain the observed column densities/profiles of this wide range is non-equilibrium processes. For example, from an apparent correlation between $N(\text{O VI})$ and $b(\text{O VI})$ Heckman et al. (2002) propose that O VI arises in post-shocked, radiatively cooling gas. Using cooling functions from Sutherland & Dopita (1993), they predict column densities of other ions that exist in both hotter gas (like that producing Ne VIII) that would exist prior to O VI, and cooler gas (like that producing O V) that exists after. In Table 4, we report the observed Ne VIII/O VI, and O V/O VI column density ratio for the broad components on the blue and red wings. In addition, we report the predicted column density ratios from Heckman et al. (2002) for gas behind a 600 km s^{-1} shock cooling from $T = 10^6 \text{ K}$ and a post-shock velocity of 100 km s^{-1} in two extremes, isobaric and isochoric cooling. From the table, only the Ne VIII/O VI column density ratio on the red wing could be explained through a radiative cooling scenario somewhere between these two extremes. However, in both components, the observed O V/O VI ratio is too large for either extreme; there is too much O V (relative to O VI) in the observed profile.

Another alternative non-equilibrium process is photoionization by a local radiation field produced by a fast,

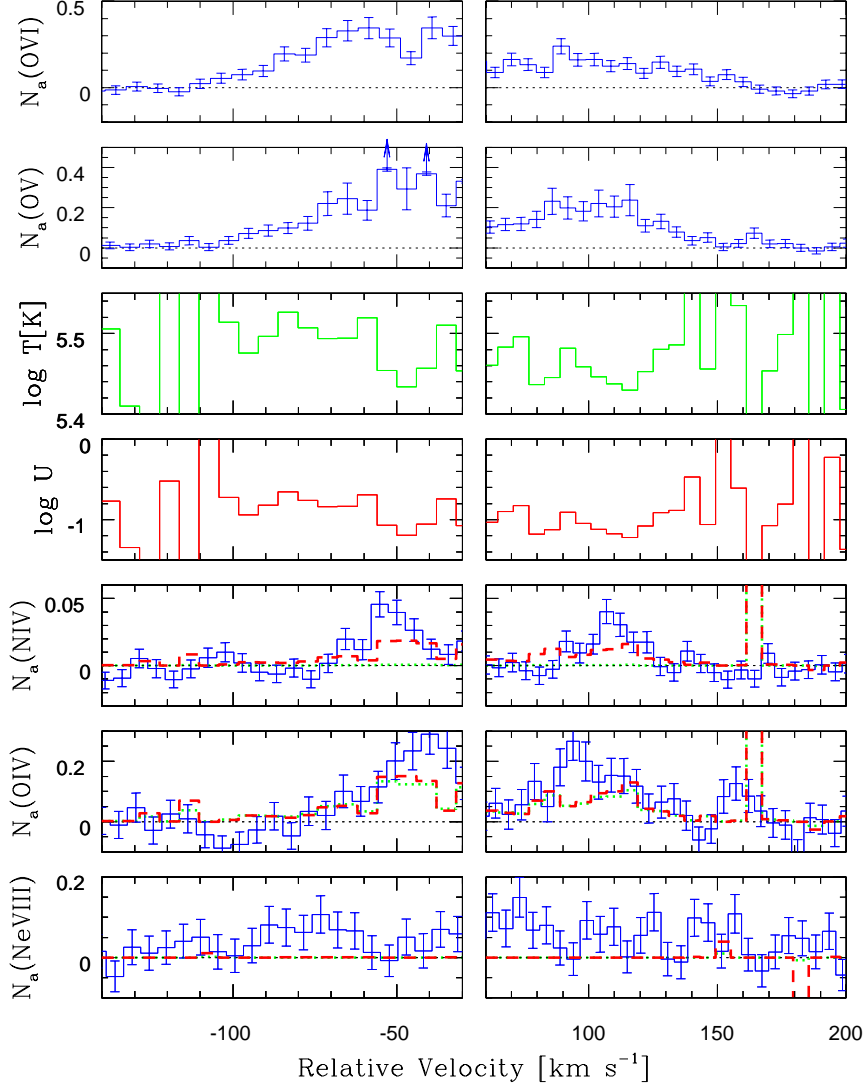


FIG. 13.— In the above panels, we show the predictions from collisional ionization equilibrium (CIE) and photoionization equilibrium (PIE) models on the apparent column density profiles of N IV, O IV, and Ne VIII, where the ionization level is tuned to match the O V–VI profiles. On the left we show the analysis of the broad component on the blue wing of the absorption-line system. On the right, we show the same for the component on the red wing. The top two panels on each side show the O VI, and O V apparent column density profiles. (Units on all apparent column density profiles are $10^{13} \text{ cm}^{-2} (\text{km s}^{-1})^{-1}$.) The next two panels show to the optimal temperature (for CIE) and ionization parameter (for PIE) as a function of velocity that reproduces the O VI/O V column density ratio. In the bottom three panels, we show the observed apparent column density profiles of N IV, O IV, and Ne VIII (solid histogram with error bars), and the predicted profiles from CIE (dotted histogram) and PIE (dashed histogram) models and assuming solar relative abundances ($\text{N}/\text{O}=0.132$, $\text{Ne}/\text{O}=0.151$).

TABLE 4. COMPARISON OF HIGH-IONIZATION SPECIES WITH NONEQUILIBRIUM MODELS

	$N(\text{Ne VIII})/N(\text{O VI})$	$N(\text{O V})/N(\text{O VI})$
$-140 \leq v [\text{km s}^{-1}] \leq -30$	0.23 ± 0.06	0.83 ± 0.12
$+60 \leq v [\text{km s}^{-1}] \leq +200$	0.58 ± 0.11	1.17 ± 0.13
Isobaric Cooling ^a	1.26	0.5
Isochoric Cooling ^a	0.5	0.6
Shock Ionization ^b	0.005–1.026	0.097–0.433

NOTE. — Errors on all quantities are quoted at 1σ confidence.

^aFrom Heckman et al. (2002), these ratios are based on radiative cooling models from Sutherland & Dopita (1993), assuming a shock velocity of 600 km s^{-1} , a post-shock velocity of 100 km s^{-1} , and an initial temperature of 10^6 K .

^bFrom Dopita & Sutherland (1996), these ratios assume shock velocities in the range $100\text{--}500 \text{ km s}^{-1}$, and magnetic parameters in the range $0\text{--}4 \mu\text{G cm}^{3/2}$.

highly radiative, magnetic pressure supported shock. Dopita & Sutherland (1996) present such a scenario and provide tables of predicted column densities of a multitude of ions for a range of shock velocities (200–500 km s^{−1}) and magnetic parameters, $B/n^{1/2}$ (0–4 μG cm^{3/2}). The range of Ne VIII/O VI and O V/O VI ratios predicted by these models is also listed in Table 4. In this case, the Ne VIII/O VI ratio of both components could be explained under this scenario. However, as with radiative cooling scenario, the predicted O V column relative to O VI is much smaller (by about a factor of two) than observed. Other non-equilibrium scenarios include turbulent mixing of entrained gas (e.g., Slavin, Shull, & Begelman 1993), and magnetic conduction through the interface between gases (e.g., Borkowski, Balbus, & Frstrom 1990). However, these papers make predictions on the column density of more commonly detected ions, such as C IV, N V, and Si IV which are not covered by our data. Thus we cannot address the applicability/validity of these scenarios to the gas in the broad components.

5. DISCUSSION

We have analyzed the detailed ionization and geometric constraints and physical conditions provided by the wide range of ionization species detected in the absorption line system associated with the quasar HE0226-4110. The rest-frame wavelength coverage of our FUSE and HST/STIS data, 610–1150 Å, provide coverage of six ionization stages of oxygen. Four stages are detected (O III–VI) and these show a remarkable ionization-dependent complexity in the gas kinematics. Comparatively low-ionization species like H I, C III, and O III are only detected in a single narrow component. Two broad components flank this component in velocity and appear mostly in high-ionization species. Intermediate-ionization stages, like N IV and O IV, are featured in both a narrow component (although with in a separate phase that is offset in velocity) and the two flanking broad components. An additional narrow component exists in this associated system that is only apparently detected in O VI. In the following sections we discuss possible origins, locations, and structural geometries for these components. [As has been our convention through-out this paper, we quote velocities relative to $z_{\text{abs}} = 0.4925$.]

5.1. Low and Moderate-ionization gas at $v = -8$ km s^{−1} (in the $z_{\text{abs}} = 0.4925$ rest-frame)

The narrow low and moderate-ionization component detected in H I, C III, and O III is the most well-constrained component in this absorption-line system. All three species appear to arise in the same phase of gas, as motivated by the kinematics of profiles. Using measurements of the H I/O III and C III/O III ratios (and limits on the O IV/O III and N IV/O III ratios), we determine the optimal metallicity, ionization parameter and total hydrogen column density of the O III-bearing gas (see equation 4). We now focus on the implications of these results on the location and geometry of the gas.

The ionization parameter is related to the spectral energy distribution of the quasar (L_ν), the distance between the absorber and the quasar central engine (r),

and the density of the absorber (n) via:

$$U = \frac{1}{4\pi r^2 n c} \int_{\nu_{\text{LL}}}^{\infty} \frac{L_\nu}{h\nu} d\nu, \quad (10)$$

where ν_{LL} is the frequency of the Lyman limit. For the spectral energy distribution specified in the photoionization models, this relationship reduces to:

$$\begin{aligned} \log U &= -1.43 - \log n_{3.5} - 2 \log r_{21.5} \\ n_{3.5} &= n / (10^{3.5} \text{ cm}^{-3}) \\ r_{21.5} &= r / (10^{21.5} \text{ cm}) = r / (1 \text{ kpc}). \end{aligned} \quad (11)$$

The reference values for the density and distance are taken from current limits for intrinsic narrow absorption lines based on variability analyses (e.g., Wise et al. 2004; Narayanan et al. 2004; Ganguly et al. 2001; Hamann et al. 1997).

Likewise, the thickness scale of the absorbing gas is given by the ratio between the total hydrogen column density and the density of the gas. Combining equations 4 and 11, we find the following relationships for the location (r) and thickness (Δr) of the absorbing gas:

$$\begin{aligned} r &= 1.5 n_{3.5}^{-1/2} \text{ kpc} \\ \Delta r &= 7.3 n_{3.5}^{-1} \text{ AU}. \end{aligned} \quad (12)$$

Before considering possible locations for the gas and related implications, we make one observation regarding the spatial extent (i.e., perpendicular to the line of sight) of the absorbing structure. Given that the absorber fully occults the UV continuum as evidenced by the zero flux levels of the C III $\lambda 977.020$ and H I Lyman β lines, we can place a lower limit on extent of the absorber on the plane of the sky. Using our fit to the optical spectrum (in particular, our characterization of the H β emission line and the underlying power law continuum), we estimate that the black hole powering the quasar has a mass of $10^{8.7-9.6} M_\odot$. The range comes from using the scaling relations of Kaspi et al. (2000) and Peterson et al. (2004). Black holes with masses in that range have a gravitational radius of $r_g = 4.6\text{--}38$ AU. The UV continuum is typically emitted in the region between $6r_g$ and $60r_g$ (e.g., Murray et al. 1995), with the bounds arising from the radius of last stable orbit (for a non-rotating black hole) and the radius at which the accretion-disk becomes optically thick. Thus the absolute minimum spatial extent of the cloud is $w \approx 600$ AU ($2 \times 60r_g$). Combining this with equation 12 The aspect ratio of the O III-bearing gas (width/thickness) is then:

$$\begin{aligned} w/\Delta r &\gtrsim 82 n_{3.5} \\ w/\Delta r &\gtrsim 82 \left(\frac{r}{1.5 \text{ kpc}} \right)^{-2} \end{aligned} \quad (13)$$

In the second expression, we have used the relationship between the source-absorber distance, and the the absorber density.

From a consideration of the statistical frequency of high-ionization (C IV, N IV, and O VI) associated absorbers in low-redshift ($z \lesssim 1$) quasars as a function of quasar redshift, Ganguly et al. (2001) placed high-ionization AAL gas in the shearing region between the outflowing from the accretion disk, and a low-density, very highly ionized medium above the wind (see their Figure 13). This scenario was physically motivated

by the hydrodynamic simulations of Proga et al. (2000) which show instabilities in that region. For the O III-bearing gas to reside in that region (within a parsec of the UV continuum source), the gas must have a density in excess of $7 \times 10^9 \text{ cm}^{-2}$, and an aspect ratio larger than 1.8×10^8 . While the density is fairly reminiscent of the broad line regions of quasars, the aspect ratio is unphysically large. Similarly, the placement of intrinsic narrow absorption lines by Elvis (2000) at the bottom of the outflow would require even larger densities and aspect ratios. This does not necessarily rule out either model, since those models only apply to high ionization absorbers. Rather, this AAL, and in particular the O III-bearing gas, points to an incompleteness in the models purporting to place gas that is related (at least statistically) to quasars. In the following subsections, we entertain two possible, alternative locations for the gas producing the O III-bearing gas, the narrow emission line region of the quasar, and the halo of the quasar host galaxy.

5.1.1. Narrow Emission Line Gas?

In principle, combining information from [O III] $\lambda 5007$ emission and O III $\lambda 832.927$ absorption, we can potentially constrain the density (n) of the absorbing gas. Density information is important since it enables an assessment of the geometry of the absorbing material courtesy of equation 9. In some associated absorbers, density constraints (e.g., from excited fine-structure lines) combined with photoionization models have indicated large distances between the associated absorber and the QSO central engine (e.g., Morris et al. 1986; Tripp et al. 1996; Hamann et al. 2001; Eracleous et al. 2003; Gabel et al. 2004). In other cases, similar analyses place the AALs close to the central engine (e.g., Srianand & Petitjean 2000). The density estimates on these rely on excited states of comparatively low-ionization species, namely C II*, Si II*, Fe II*, and C III* that are not widely detected in AALs, so it is not clear how broadly applicable the results of those studies are in understanding the distribution of location of AAL gas. The potential association of O III absorption in the narrow-emission line region, then, has paramount importance to this issue, since strong narrow emission lines are far more common in AGN.

In trying to associate this absorption-line gas with the narrow emission line gas observed in [O III] $\lambda 5007$, it is important to consider the geometry and kinematics of the emitting region compared to the pencil-beam probe producing the absorption. We note a few basic facts (in addition to those mentioned above regarding physical conditions and geometry) that a narrow emission line region model incorporating the O III-absorbing gas should address: (1) a match in velocity ($z_{\text{em}}([\text{O III}]) = 0.4928 \pm 0.0001$, $z_{\text{abs}}(\text{O III}) = 0.49246 \pm 0.00006$, $\Delta v = 68 \pm 23 \text{ km s}^{-1}$; see Table 1 and §4) between the absorption and emission line centroid; (2) a factor of 32 difference in the velocity widths (FWHM $\sim 470 \text{ km s}^{-1}$ for emission, FWHM $\sim 15 \text{ km s}^{-1}$ for absorption); and (3) an [O III] $\lambda 5007$ luminosity of $3.2 \times 10^{42} \text{ erg s}^{-1}$.

We first consider whether gas with $\log U = -2.29$ and $[\text{O}/\text{H}] = +0.12$ (as implied by the analysis of the absorption lines) can reproduce the observed luminosity. To do so, we ran another grid of CLOUDY models over density, placing the gas at a distance that reproduces an ioniza-

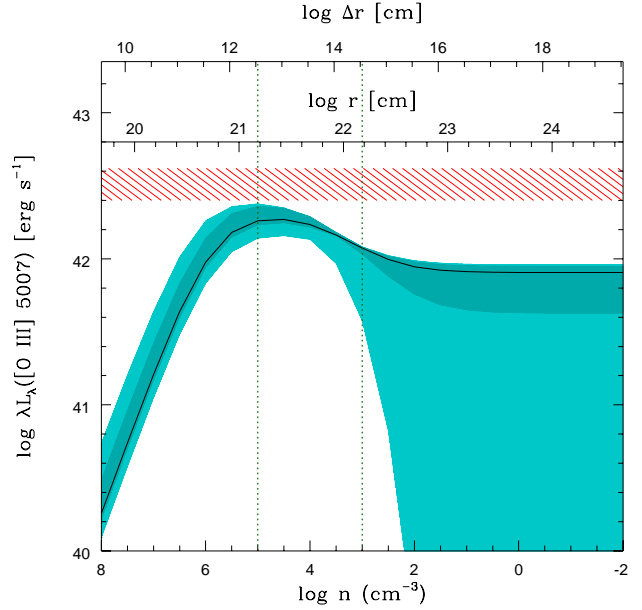


FIG. 14.— In the above figure, we plot the predicted [O III] $\lambda 5007$ line luminosity as a function of gas density. The solid curve shows the grid of CLOUDY models using $\log U = -2.29$, and $[\text{O}/\text{H}] = +0.12$. The shaded regions around the solid curve represent the set of models allowed by the 1σ and 2σ uncertainties in the ionization parameter and metallicity from the absorption-line diagnostics (see Figures 7-8). The hatched region at $\log \lambda L_{\lambda}([\text{O III}] \lambda 5007) [\text{erg s}^{-1}] = 42.4-42.62$ marks the 1σ confidence luminosity of the EM6 emission line component which we associate with this gas. The dashed vertical lines at $\log n [\text{cm}^{-3}] = 3$ and 5 indicate the typical range of densities expected for the narrow emission line region gas using emission-line diagnostics. The top two axes indicate the inferred absorber thickness (Δr , assuming a total hydrogen column density $\log N(\text{H}) [\text{cm}^{-2}] = 17.54$) and distance from the accretion disk (r). These axes are implicitly dependent on the ionization parameter, and therefore only apply to the solid curve.

tion parameter of $\log U = -2.29$. At each grid point, we compute the [O III] $\lambda 5007$ line luminosity and compare that to the observed line luminosity for emission-line component 6 from our fit to the [O III] $\lambda 5007$ line in §3 (see Table 1), since this is the component whose redshift matches that of the O III $\lambda 832.927$ absorption. In these models, the gas is assumed to exist in the thin spherical shell that surrounds the point-like central engine. The absorption arises from the pencil-beam probe through the shell toward the observer, while the emission arises (isotropically) from the integration over the entire shell.

The results of the grid of CLOUDY models are shown in Figure 14. On the top axes of the plot, we show the implied distance, $\log r [\text{cm}] = 21.5 + 0.5(4.36 - \log n [\text{cm}^{-3}])$, and thickness, $\log \Delta r [\text{cm}] = 17.54 - \log n [\text{cm}^{-3}]$. The model curve shows a power-law decrease of the [O III] $\lambda 5007$ line luminosity with increasing density starting at $\log n [\text{cm}^{-3}] \sim 5$. At smaller densities, the curve falls off and flattens to a constant line luminosity whose value depends on the ionization parameter and metallicity. The shaded regions around the plotted curve represent families of models that lie in the 1σ and 2σ confidence regions shown in Figure 8. Models with higher ionization parameters and/or metallicities decrease the constant luminosity to which the curve levels out. The hatched region across the top of the figure is vertically placed at the

observed luminosity of the EM6 component, with the thickness of the region representing the 1σ confidence interval. The vertical dotted lines are placed to indicate the range of densities, $\log n[\text{cm}^{-3}] = 3 - 5$, commonly associated with narrow emission line region gas. The plot shows that there is no density at which the model predicts sufficient line luminosity to match the observation, although at the $\sim 1.5\sigma$ level, there is reasonable agreement around $n \sim 10^5 \text{ cm}^{-3}$.

In principle, the discrepancy could be further explained by the fragmentation of the shell into several smaller clouds within the narrow emission line region (with only one cloud producing the absorption) or with the inclusion of other important physical processes (such as shocks). Such an explanation may be required in order to reconcile the model with the observed velocity width of the EM6 component ($\sim 470 \text{ km s}^{-1}$). A static thin shell with an internal velocity dispersion of $\sim 15 \text{ km s}^{-1}$ (as indicated by the pencil-beam absorption-line probe) does not yield an overall velocity width of 470 km s^{-1} from the emission-line gas. [In order to explain the 68 km s^{-1} blueshift of the absorption line relative to the emission line, one might invoke an expanding thin shell, but even such a shell would not produce the observed emission line velocity width.] With a fragmented shell, a possible solution is that the bulk velocity dispersion of the resulting clouds (arising perhaps from orbital motion about the central black hole) is several hundred km s^{-1} , while the internal velocity dispersion of any single cloud is as indicated by the absorption-line profiles. This scenario has two advantages over a simple thin shell model: (1) a thin shell may have stability problems, and (2) the added surface area from fragmentation could yield the factor of two difference between the predicted and observed [O III] $\lambda 5007$ luminosities.

However, there are other potential self-consistency problems with the absorbing gas with a density of $n \sim 10^5 \text{ cm}^{-3}$. The implied distance between the absorber and the accretion disk is $r \approx 0.5n_5^{-1/2} \text{ kpc}$. While this is reasonable for a location within the narrow emission line region (e.g. Schmitt et al. 2003; Bennert et al. 2002), the implied cloud thickness is extremely small, $\Delta r \approx 0.24n_5^{-1} \text{ AU}$, implying a minimum aspect ratio (see equation 13) of $\approx 1200n_5$ for the cloud. For the model to provide a self-consistent representation of data, the gas probed by the absorption must be distributed in the form of a sheet or a closed spherical shell (or sets of sheets/shells), rather than a discrete almost-spherical cloud. Possibly, such a deformation might result from the hydrodynamics of gas orbiting a central supermassive black hole, although such a deformed cloud would probably fragment further. Another potential problem with this scenario is that only one cloud is seen in absorption. While it is possible that only one cloud would be intercepted by the sightline in this patchy cloud scenario, it is improbable. Only future observations to look for variability in the O III-absorption profile and/or higher resolution observations (spatial and spectral) of the [O III] $\lambda 5007$ can address the patchiness of the narrow emission line region and the number of clouds intercepted by the absorption-line probe.

A number of additional observational tests are required to fully explore the implications of this result. Additional

data covering the O III $\lambda 832.927$ line in the associated absorption of other quasars is required to ascertain the frequency in which gas in the narrow emission line region is observed in absorption and the fraction of associated absorbers that originate in this region. Since this associated absorber clearly has many components with different gas phases, it is likely that several different regions contribute to the observed profiles. Spectra with high-resolution ($\lambda/\Delta\lambda \geq 10^4$) are required to disentangle the relative importance of the different quasar emitting regions to the observed frequency of associated absorbers. For this and other quasars whose spectra feature associated absorption in O III $\lambda 832.927$, higher-resolution spectroscopy of the [O III] $\lambda 5007$ line would be useful in directly, and more precisely, comparing the kinematics of the extended narrow emission-line region with the pencil-beam probe provided by the O III absorption. The density diagnostics available from the analysis of emission lines would greatly assist in disentangling the geometry of intrinsic absorbers.

5.1.2. Quasar Host Galaxy Halo Gas?

Another potential location for the O III-absorbing gas is the halo of the quasar host galaxy. If the gas were located, for example, 10 kpc away from the central engine, then (by equation 12) the density of the gas would be $\sim 230 \text{ cm}^{-3}$, and the thickness of the gas would be $\sim 100 \text{ AU}$. The resulting aspect ratio (by equation 13) would be > 6 , so such a structure would still be rather non-spherical. Are there local analogues of such a structure? High-velocity clouds such as Complex C or the Magellanic Stream may have such aspect ratios, but do not have such large densities (e.g., Sembach et al. 2001; Wakker 2001; Tripp et al. 2003; Fox et al. 2004, 2005; Wakker 2004). If the gas were $\gtrsim 13.5 \text{ kpc}$ away from the central engine, also within the halo of the quasar host galaxy, that would allow for a spherical structure with an absorbing thickness scale $> 600 \text{ AU}$. At such a distance, the density is smaller than 37 cm^{-3} . However, clouds in the halo (or outer corona) of the Galaxy do not have the high (super-solar) metallicity (e.g., Wakker 2001; Richter et al. 2001; Tripp et al. 2003; Wakker 2004) observed in the O III-absorbing gas.

While there are no local analogues of an absorbing structure that exists in the halo with super-solar metallicity and a large density, such an idea is not without precedent. The associated absorber observed toward 3C 191 (Hamann et al. 2001) features very similar properties - gas with a density of $\sim 300 \text{ cm}^{-3}$, residing 28 kpc from the quasar central engine. However, that absorber is detected in very low-ionization gas (Mg II, C II, Si II) and excited state lines from C II* and Si II*, which are not present in the O III-absorbing gas in this associated absorber. Moreover, the low-ionization lines observed in the 3C 191 associated absorber exhibit the signature of partial coverage (dilution of the absorption line troughs unocculted light) and the coverage fractions imply gas clouds with a spatial extent smaller than 0.01 pc . (The thickness of the 3C 191 associated absorber is unknown as Hamann et al. (2001) are only able to place upper limits on the total hydrogen column density.) In addition, the 3C 191 associated absorber is apparently outflowing (it is blueshifted by more than 400 km s^{-1} relative to the broad emission lines), whereas the O III-absorbing

gas in this associated absorber is at a velocity fairly coincident with the [O III] $\lambda 5007$ narrow emission line. It is possible that O III gas in this associated absorber is a slightly higher-ionization analogue of that absorber. If so, then the speculation by Hamann et al. (2001) that the gas may be residue from a nuclear starburst superwind may be applicable. It is unlikely that this material is the result of an ejection from a BAL wind as the gas appears to be at rest (or close to it) with respect to the narrow emission region, which would imply a severe deceleration (by several order of magnitude).

5.2. *Moderate and High-ionization gas at $v=0$ \mathcal{E} $+50 \text{ km s}^{-1}$ (in the $z_{\text{abs}} = 0.4925$ rest-frame)*

Additional narrow features are detected in this associated absorption-line system, but the physical conditions of the gas are much less constrained than the O III-bearing gas. Both of these components have a higher ionization than the other narrow component. The $v = 0 \text{ km s}^{-1}$ component is detected as a single narrow feature in intermediate-ionization species such as N IV, O IV, and S IV, while the $v = +50 \text{ km s}^{-1}$ component is detected as a narrow feature in O VI only. Absorption from other intermediate- and high-ionization species - O V and Ne VIII - is evident, but kinematics of these species do not trace the narrow components. There is no detectable H I that is associated with either of these components, as all the observed H I appears over velocity range and has the kinematic appearance of the O III-bearing gas. Whether there is a relationship between these components either to each other, or to the O III-bearing gas is unclear. The gas could lie close to the quasar central engine, or far out in a companion galaxy.

The velocity widths of these two components are fairly narrow [$b(v = 0 \text{ km s}^{-1}) \sim 15 \text{ km s}^{-1}$, $b(v = +50 \text{ km s}^{-1}) \sim 9 \text{ km s}^{-1}$]⁷ and, in the case of the $v = +50 \text{ km s}^{-1}$ component, suggest that collisional processes are not important in the ionization, leaving radiative processes as the likely dominant source of ionization. Even if the gas producing these components were far from the central engine, perhaps in a dwarf companion to the quasar host galaxy, photoionization by the central engine is likely the dominant source of ionizing photons. For the $v = +50 \text{ km s}^{-1}$ narrow component, we used the available constraints from O V and Ne VIII to infer that the gas must have an ionization parameter in the range $-0.35 \lesssim \log U \lesssim -0.02$. However, since there is no information on the density of this component (e.g., via association with a line-emitting region, or excited state transitions, or time-variability), it is difficult to know where the gas is located relative to quasar central engine [since the distance between the accretion disk and the absorber is degenerate with the density ($U \propto L/nr^2$)].

One possible interpretation is that these components arise in structures similar to that of the other narrow component, but with smaller densities (thus producing higher-ionization species). For the narrow O VI component, if the gas lies within the narrow emission line region of the quasar, then it has a density in the range $10^{2.69} \lesssim n[\text{cm}^{-3}] \lesssim 10^{3.02}$. This is consistent with the range of densities typically found in quasar nar-

row emission line regions. The gas producing the narrow O VI absorption would not contribute significantly to the formation of the narrow emission lines since the gas is too ionized. If the total hydrogen column density of the narrow O VI component is in the range $17.75 \lesssim \log N(\text{H})[\text{cm}^{-2}] \lesssim 18.1$ (as would be implied if the metallicity were similar to the O III component), then the thickness of the gas sheet producing the narrow O VI component lies in the range $35 \lesssim \Delta r[\text{AU}] \lesssim 170$, and an aspect ratio of at least 3.5. Using the absolute maximum total hydrogen column density (as derived by the non-detection of H I and the limiting ionization fraction), the aspect ratio cannot be smaller than 1.4 (if the gas lies with the narrow emission line region).

Alternatively, if the narrow O VI gas is in the halo of the quasar host galaxy, 10kpc away from the central engine, then it would have a density in the range $0.5 \lesssim n[\text{cm}^{-3}] \lesssim 1.1$, which would be consistent with local versions of halo gas clouds. While these are only two of many possible scenarios for these higher-ionization components, they have the attraction of readily explaining why all three narrow components arise within the same absorption-line system very close in velocity (within 50 km s^{-1}).

It is noteworthy that Bergeron et al. (2002) and Carswell, Schaye, & Kim (2002) detect O VI components at $z_{\text{abs}} \gtrsim 2$ which are also narrow enough to imply photoionization. In particular, Bergeron et al. (2002) detect a component at $z_{\text{abs}} = 2.36385$ in the direction of Q 0329-385 ($z_{\text{em}} = 2.423$) with a similar b -value as the $v = +50 \text{ km s}^{-1}$ component in this associated system. [The difference in redshift corresponds to a velocity separation of 5225 km s^{-1} , which is close to the arbitrary cutoff for associated absorption-line systems.] That component is detected in H I and has a $N(\text{O VI})/N(\text{H I})$ ratio of ≈ 0.09 . Our narrow O VI component is not detected in H I and has a ratio of $N(\text{O VI})/N(\text{H I}) \gtrsim 5$. In fact, in all of the cases presented by both Bergeron et al. (2002) and Carswell et al. (2002), the narrow, hence photoionized, O VI is accompanied by detection of H I, in marked contrast to the $v = +50 \text{ km s}^{-1}$ component in this associated absorption-line system. [Carswell et al. (2002) note that in all cases that they present, the H I column density is larger than $2 \times 10^{14} \text{ cm}^{-2}$. None of their components appear to have metallicities in excess of $[\text{C}/\text{H}] \approx -1.7$.] Bergeron et al. (2002) report that their photoionized-O VI components (with only one exception) have metallicities in range $-2.5 \leq [\text{C}/\text{H}] \leq -0.5$, with super-solar $[\text{O}/\text{C}]$. This is consistent with our constraint on $[\text{O}/\text{H}]$ (≥ -0.62) for the $v = +50 \text{ km s}^{-1}$ component.

In many of the cases studied by Bergeron et al. (2002) (including the one mentioned above), the narrow O VI components are accompanied by detections of C IV (and occasionally N V). For the narrow O VI component in this absorption-line system, the strong UV doublet from neither species is covered by our spectra. If the photoionization models described in §4.1.3 are an accurate representation of the O VI-bearing gas, then we predict column densities in the range $12.21 \lesssim \log N(\text{C IV}) \lesssim 12.4$, and $12.74 \lesssim \log N(\text{N V}) \lesssim 12.82$ (for $[\text{O}/\text{C}] = [\text{O}/\text{N}] = 0$). For comparison, the associated component from the Bergeron et al. (2002) study mentioned above has $\log N(\text{C IV}) = 12.33$, and $\log N(\text{N V}) = 12.00$. The C IV column density is within the range predicted

⁷ Both of these values were obtained from direction integration of the second moment of the apparent optical depth profile.

for the $v = +50 \text{ km s}^{-1}$ component (even if $[\text{O}/\text{C}]$ were slightly super-solar), but the N V column density is smaller than that of the $v = +50 \text{ km s}^{-1}$ component (if N/O is solar). If the N/O abundance in the $v = +50 \text{ km s}^{-1}$ component is subsolar by 0.7 dex, then the two components are very similar aside from the differences in redshift and H I column density. Bergeron et al. (2002) find precisely such an abundance requirement for their $z_{\text{abs}} = 2.25146$ component ($[\text{O}/\text{N}] \geq 0.7$). If this is a general property, then the $v = +50 \text{ km s}^{-1}$ component in this associated absorption-line system is very consistent with the populations of photoionized-O VI discussed by Bergeron et al. (2002) and may be higher metallicity analogues to those studied by Carswell et al. (2002).

5.3. Broad Components

Intermediate and high-ionization species, namely N IV, O V-VI, and Ne VIII, show absorption in two relatively broad components. In O VI, the absorption in these two components constitutes roughly 30% of the total integrated column density. The kinematics of O V and O VI appear to trace each other, suggesting that the two arise in the same phase of gas. Moreover, there is no evidence for trough-dilution by unocculted quasar light, so it is unlikely that these components arise from the quasar outflow (depending on the orientation of the accretion disk and possible presence of a scattering medium).

Since many different physical process could be involved in the production of this gas (photoionization by the quasar central engine, collisional ionization from shocks, etc.), it is unclear what ionization model to apply in order to infer physical conditions. In §4.2, we explored various ionization models, collisional ionization equilibrium (CIE), photoionization equilibrium (PIE), radiative cooling (RC) as describe by Heckman et al. (2002), and shock ionization (SI) as described by Dopita & Sutherland (1996). CIE and PIE models that are tuned (through the temperature or ionization parameter) to reproduce the observe O V and O VI column densities are unable to fully explain the presence of N IV, O IV, and Ne VIII in the same phase of gas (Fig. 13). The RC and SI models are unable to even simultaneously reproduce the O V and O VI column densities under the range of parameters considered (see Table 4).

Of the models considered, photoionization by the quasar is able to explain the largest number of species simultaneously. However, such a model requires additional phases to explain the presence of Ne VIII, and some of the N IV and O IV (and it is not clear what ionization models are appropriate for such phases). For the phase producing O V-VI, if photoionization by the quasar is the correct model, then O VII should be the dominant ionization stage in these components. An X-ray spectrum (e.g. with *Chandra* or *XMM-Newton*) to look for O VII absorption would help to elucidate the important physical processes producing the broad components. Although the total column density of O VII predicted by these components ($N \lesssim 10^{15} \text{ cm}^{-2}$) would probably not be detectable, the higher ionization gas producing the Ne VIII may be detectable. Information on the higher ionization stages of oxygen would compete the picture of the ionization structure of gas in this absorption-line system.

Like the narrow higher-ionization components, there is

no density information in the broad components. Thus, it is not possible to transform the optimal ionization parameter into a distance. Likewise, there is no metallicity information, so it is not possible to robustly gauge the thickness of the absorbing regions. There are several plausible locations for this gas (ablation off an obscuring torus, a diffuse medium co-spatial with the narrow emission-line region, satellite dwarf galaxies, etc.) and we cannot distinguish between them.

6. SUMMARY OF RESULTS

We have obtained high-resolution, rest-frame extreme ultraviolet spectra of the quasar HE 0226-4110. The spectra cover the wavelength range 610–1150 Å in the rest-frame of the quasar. We detect an associated absorption line system in a wide range of ionization species, including several Lyman series lines, Ne VIII, and four adjacent stages of oxygen, O III-VI. Strong transitions from O I and O II are also covered by our data, but are not detected. The high quality of these spectra allow us to study the kinematic structure of the gas. A summary of our results follows.

1. The O III $\lambda 832.927$ line is detected in a single narrow component at $v = -8 \text{ km s}^{-1}$ in the $z = 0.4925$ reference frame. The kinematics of the absorber indicate that all the detected H I and C III are associated with this component, and we use the O III profile to gauge the column density ratios of the narrow emission line region gas. Using CLOUDY photoionization models to constrain physical parameters, we find that this cloud has a metallicity of $[\text{O}/\text{H}] = +0.12^{+0.16}_{-0.03}$, an ionization parameter $\log U = -2.29^{+0.02}_{-0.23}$, and a total hydrogen column density of $\log N(\text{H}) = 17.54^{+0.04}_{-0.25}$.
2. We detect additional narrow components at 8 km s^{-1} , and 58 km s^{-1} redward of the O III-bearing gas (at $v = 0$ and $+50 \text{ km s}^{-1}$ in the $z = 0.4925$ rest-frame, respectively). These components are detected in higher-ionization species. The $v = 0 \text{ km s}^{-1}$ component is detected as a single narrow feature in N IV, O IV, and S IV, while the $v = +50 \text{ km s}^{-1}$ component is similarly detected only in O VI. Absorption from O V and Ne VIII is detected at the same velocities, but the kinematics rule out a direct association. The O VI profile is too narrow to be produced by collisional processes. Using limits on the Ne VIII/O VI and O V/O VI column density ratios, we constrain the ionization parameter of the $v = +50 \text{ km s}^{-1}$ component to the range $-0.35 \lesssim \log U \lesssim -0.02$.
3. Two broad, smooth components are also detected in N IV, O V-VI and Ne VIII. We find no evidence for the dilution of troughs expected from gas arising in an accretion-disk outflow. There is insufficient information to propose a unique ionization scenario (both collisions and photons could contribute to the ionization). Radiative cooling and shock ionization models are unable to produce sufficient amounts of O V in the ranges of model parameters considered by Heckman et al. (2002) and Dopita & Sutherland (1996) to match the observed

column density. Photoionization by the quasar is able to explain some N IV and O IV in a phase tuned to produce O V-VI. However, other phases are required to fully explain the observed profiles. A separate phase is required, regardless of mechanism, to explain the presence of Ne VIII.

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